Mon. Not. R. Astron. Soc. **000**, 1–?? (xxxx) Printed 29 April 2010 (MN LATEX style file v2.2)

Color and stellar population gradients in galaxies. Correlation with mass.

C. Tortora[1,2⋆], N.R. Napolitano[3], V.F. Cardone[4], M. Capaccioli[2], P. Jetzer[1]
R. Molinaro[3,5]

[1] *Universität Zürich, Institut für Theoretische Physik, Winterthurerstrasse 190, CH-8057, Zürich, Switzerland*
[2] *Dipartimento di Scienze Fisiche, Università di Napoli Federico II, Compl. Univ. Monte S. Angelo, 80126 - Napoli, Italy*
[3] *INAF – Osservatorio Astronomico di Capodimonte, Salita Moiariello, 16, 80131 - Napoli, Italy*
[4] *Dipartimento di Fisica Generale A. Avogadro, Università di Torino and Istituto Nazionale di Fisica Nucleare - Sezione di Torino, Via Pietro Giuria 1, 10125 - Torino, Italy*
[5] *Dipartimento di Fisica, Politecnico di Torino, Corso Duca degli Abruzzi 24, 10129 - Torino, Italy*

Accepted Received

ABSTRACT

We analyze the color gradients (CGs) of $\sim$ 50 000 nearby Sloan Digital Sky Survey (SDSS) galaxies estimated by their photometrical parameters (Sérsic index, total magnitude, and effective radius). From synthetic spectral models based on a simplified star formation recipe, we derive the mean spectral properties, and explain the observed radial trends of the color as gradients of the stellar population age and metallicity. Color gradients have been correlated with color, luminosity, size, velocity dispersion, and stellar mass. Distinct behaviours are found for early- and late-type galaxies (ETGs and LTGs), pointing to slightly different physical processes at work in different morphological types and at different mass scales.

In particular, the most massive ETGs ($M_* \gtrsim 10^{11}\,M_\odot$) have shallow (even flat) CGs in correspondence of shallow (negative) metallicity gradients. In the stellar mass range $10^{10.3-10.5} \lesssim M_* \lesssim 10^{11}\,M_\odot$, the metallicity gradients reach their minimum of $\sim -0.5\,\mathrm{dex}^{-1}$. At $M_* \sim 10^{10.3-10.5}\,M_\odot$, color and metallicity gradient slopes suddenly change. They turn out to anti-correlate with the mass, becoming highly positive at the very low masses, the transition from negative to positive occurring at $M_* \sim 10^{9-9.5}\,M_\odot$. These correlations are mirrored by similar trends of CGs with the effective radius and the velocity dispersion. We have also found that age gradients anti-correlate with metallicity gradients, as predicted by hierarchical cosmological simulations for ETGs. On the other side, LTGs have gradients which systematically decrease with mass (and are always more negative than in ETGs), consistently with the expectation from gas infall and supernovae feedback scenarios.

Metallicity is found to be the main driver of the trend of color gradients, especially for LTGs, but age gradients are not negligible and seem to play a significant role too. Owing to the large dataset, we have been able to highlight that older galaxies have systematically shallower age and metallicity gradients than younger ones.

The emerging picture is qualitatively consistent with the predictions from hydrodynamical and chemo-dynamical simulations. In particular, our results for high-mass galaxies are in perfect agreement with predictions based on the merging scenario, while the evolution of LTGs and younger and less massive ETGs seems to be mainly driven by infall and SN feedback.

Key words: dark matter – galaxies : evolution – galaxies : galaxies : general – galaxies : elliptical and lenticular, cD.

1 INTRODUCTION

Galaxy formation is a complicated matter which has not yet come to a complete and coherent understanding. The

⋆ E-mail: ctortora@na.astro.it

standard cosmological paradigm, the so-called ΛCDM, predicts that dark matter (DM) haloes evolve hierarchically since early epochs, with smaller units merging into massive structures (Kauffmann et al. 1993, de Lucia et al. 2006, Ruszkowski & Springel 2009), which find positive evidences in the strong size evolution in massive galaxies since $z \sim 3$ (Glazebrook et al. 2004, Daddi et al. 2005, Trujillo et al. 2006). This scenario is at variance with the evidences that today's high mass galaxies formed most of their stars in earlier epochs and over a shorter time interval than low-mass ones. This "downsizing" scheme (Cowie et al. 1996; Gallazzi et al. 2005; Jimenez et al. 2005; Nelan et al. 2005; Thomas et al. 2005; Treu et al. 2005; Bundy et al. 2006; Cimatti et al. 2006; Pannella et al. 2006; Panter et al. 2007; Fontanot et al. 2009) seems at odd with the ΛCDM hierarchical growth and seems rather to support a "monolithic–like" formation scheme (see e.g. Chiosi & Carraro 2002) which broadly predicts an inside-out formation of stars after an early dissipative collapse.

However, it is increasingly clearer that the downsizing is compatible with the hierarchical model if the feedback processes in the mass accretion history of the ΛCDM dark haloes are taken into account (see e.g., Neistein et al. 2006, Cattaneo et al. 2008, Conroy & Wechsler 2008, Cattaneo et al. 2010). On the one hand, the gas cooling and the shock heating (Dekel & Birnboim 2006, Cattaneo et al. 2008) are the main drivers of the star formation (SF) activity during the hierarchical growth of the DM haloes; on the other side supernovae (SNe), active galactic nuclei (AGN), merging, harassment, strangulation etc., generally inhibit the stellar formation (Dekel & Silk 1986; Recchi et al. 2001; Pipino et al. 2002; Scannapieco et al. 2006; Di Matteo et al. 2005; de Lucia et al. 2006; Cattaneo et al. 2006; Kaviraj et al. 2007; Schawinski et al. 2007; Cattaneo et al. 2008; Antonuccio-Delogu & Silk 2008; Khalatyan et al. 2008; Romeo et al. 2008; Tortora et al. 2009b; Weinmann et al. 2009).

The different processes might rule the star formation at the global galaxy scale, or act at sub-galactic scales (e.g. the nuclear regions vs outskirts) such that they are expected to introduce a gradient of the main stellar properties with the radius that shall leave observational signatures in galaxy colours.

This paper is motivated by the fact that color gradients (CGs) are efficient markers of the stellar properties variations within galaxies, in particular as they mirror the gradients of star ages and metallicities (Sandage 1972), although it is not yet fully clear whether metallicity is the main driver of the CGs in normal spheroidal (Saglia et al. 2000; Tamura & Ohta 2000; La Barbera et al. 2002, 2003; Kobayashi 2004; Spolaor et al. 2009; Rawle et al. 2009) and late-type systems (MacArthur et al. 2004, Taylor et al. 2005), or age plays also an significant role (Saglia et al. 2000; La Barbera et al. 2003; MacArthur et al. 2004; Spolaor et al. 2009; Rawle et al. 2009).

CGs are primarily a tool to discriminate the two broad formation scenarios (monolithic vs hierarchical), but more importantly they provide a deeper insight on the different mechanisms ruling the galaxy evolution. As these mechanisms depend on the galaxy mass scale, the widest mass (and luminosity) observational baseline is needed to remark the relative effectiveness of the different physical processes and their correlation with the observed population gradients.

The observational picture is so complex to justify a schematic review of the evidences accumulated so far. The facts are as follows.

1) On average, nearby elliptical and spiral galaxies are bluer outwards (Franx, Illingworth & Heckman 1989, Peletier et al. 1990a,b, Balcells & Peletier 1994, Tamura & Ohta 2003), while dwarfs show mainly redder outskirts (Vader et al. 1988, Kormendy & Djorgovski 1989, Chaboyer 1994, Tully et al. 1996, Jansen et al. 2000). Recently, deeper investigations of later-type galaxies have been done and have highlighted the presence of non-monotonic colour and stellar population gradients in these systems (MacArthur et al. 2004, Bakos, Trujillo & Pohlen 2008, MacArthur et al. 2009, Martínez-Serrano et al. 2009, Sánchez-Blázquez et al. 2009).

2) Spectral line indices (mainly measured in early-type galaxies, ETGs) are somehow a more efficient tracer of the stellar population properties and generally found to change with the radius like CGs (Kobayashi & Arimoto 1999, Kuntschner et al. 2006, MacArthur et al. 2009, Rawle et al. 2009).

3) Interpreting CGs in terms of metallicity gradients, in high-mass galaxies typically $d\log Z/d\log R \sim -0.3$, which is shallower than predicted by simulations of dissipative "monolithic" collapse (where -0.5 or a steeper value is expected, see e.g. Larson 1974, 1975, Carlberg 1984, Arimoto & Yoshii 1987, Kobayashi 2004, see also below).

4) There are contradictory evidences of a correlations of CGs with galaxy mass and luminosity. Earlier studies have claimed a weak correlation with the physical properties of galaxies (e.g., mass, luminosity, etc., Peletier et al. 1990a, Davies et al. 1993, Kobayashi & Arimoto 1999, Tamura & Ohta 2003), at variance with the typical monolithic collapse predictions. Recently, a stronger correlation with mass has started to emerge (e.g. Forbes et al. 2005), pointing to a metallicity gradient decreasing with the mass for low-mass galaxies (Spolaor et al. 2009, Rawle et al. 2009), accordingly with the monolithic scenario. Instead, high-mass galaxies show too shallow gradients to match the predictions of the monolithic collapse, but compatible with galaxy merging.

This wealth of observational evidences must be confronted with the model predictions coming from the different galaxy formation scenarios.

1) Steep gradients are expected when stars form during strong dissipative (monolithic) collapses in deep potential wells of galaxy cores where the gas is more efficiently retained, with a consequent longer star formation activity and a longer chemical enrichment in the inner than in outer regions (negative metallicity gradients). On top of that, the delayed onset of winds from supernovae, causing a further metal supply in the central regions, would contribute to reinforce the steepness of these gradients (see, e.g. Pipino et al. 2008). As these processes are regulated by the galaxy potential depth, which is somehow related to the galaxy mass (and luminosity) [1], high-mass galaxies are ex-

[1] It remains to see whether there might be a correlation with the the DM fraction of the systems which is also a function

pected to have metallicity gradients which are steeper than lower-mass ones. The latter seem to have almost no gradients (Gibson 1997, Chiosi & Carraro 2002, Kawata & Gibson 2003), although the results here are based on limited galaxy samples.

2) Within the hierarchical picture, merging induce a meshing of stellar populations, and lead to a more uniform metallicity distribution and to shallower CGs (White 1980, Kobayashi 2004). Strong jets from powerful AGNs can also quench the star formation on the global galaxy scale and flatten colors gradients in the host systems (Tortora et al. 2009b).

3) Metallicity and color gradients seem also to depend on the efficiency of the dissipative processes in dark haloes (Hopkins et al. 2009a), with just a weak dependence on the remnant mass (Bekki & Shioya 1999, Kobayashi 2004, Hopkins et al. 2008b), in a way more similar of the observed gradients. However, while merging is crucial to smear out the color radial variation in high-mass galaxies (Kobayashi 2004), it seems unimportant for low-mass systems (de Lucia et al. 2006, Cattaneo et al. 2008). Instead, the energy from stellar winds and supernovae and the effect of dissipative collapse might induce even positive steep gradients (Mori et al. 1997).

CGs are the most direct observables to investigate the effect of the physical processes (such as merging, AGN, SNe, stellar feedback), which drive the galaxy evolution as a function of the main galaxy parameters: luminosity, mass and central velocity dispersion. In this paper we analyze optical CGs in about 50 000 SDSS local galaxies, spanning a wide range of luminosities and masses. CGs are obtained in an indirect way from the photometrical parameters of the individual galaxies (effective radius, Sérsic index, total magnitude). Using synthetic spectral models, we compute age and metallicity gradients for early- and late-type systems, and analyze the trends with mass. The sample allows us to investigate the distribution of the CGs over an unprecedented baseline of galaxy sizes, luminosities, velocity dispersions, and stellar masses. The observed trends are interpreted by means of quite different physical phenomena at the various mass scales. This approach, while exposed to the uncertainties on the structural parameters (e.g. Sérsic index, n, and effective radius, $R_{\rm eff}$) has the advantage of dealing with large statistics; furthermore it is exportable to higher redshift galaxy surveys which are generally limited to rest-frame visual bands. In this respect we pay particular care in the check of the consistency of our results with independent analyses. We will mainly concentrate on ETGs, for which a wide collection of results (having an homogenized gradient definition) is available from literature. On the other hand, due to the uncertainties of the structural parameters of late-type galaxies (LTGs), we will take the results on this galaxy sample with the right caution, being well aware that our findings might be only a benchmark for more accurate analyses.

of the galaxy stellar mass (see e.g., Napolitano et al. 2005, Cappellari et al. 2006, Conroy & Wechsler 2008, Tortora et al. 2009a).

In §2 we present the data sample and spectral models. In §3 we show the main results of our analysis, discussing CGs as a function of structural parameters, stellar masses, and stellar properties derived from the fitting procedure. In the same Section, age and metallicity gradients are shown, and discussed within the galaxy formation scenarios in §4. We finally draw some conclusions in §5.

In the following, we use a cosmological model with $(\Omega_m,\, \Omega_\Lambda,\, h) = (0.3,\, 0.7,\, 0.7)$, where $h = H_0/100\,{\rm km\,s^{-1}\,Mpc^{-1}}$ (Spergel et al. 2003, 2007), corresponding to a Universe age of $t_{\rm univ} = 13.5\,{\rm Gyr}$.

2 DATA

2.1 Galaxy sample

Our database consists of 50 000 low redshift ($0.0033 \leqslant z \leqslant 0.05$) galaxies in the NYU Value-Added Galaxy Catalog extracted from SDSS DR4 (Blanton et al. 2005b, hereafter B05)[2]. B05 have performed a set of quality checks contemplating the analysis of large and complex galaxies, incorrectly deblended in the SDSS pipeline, and of the star/galaxy separation, with a number of eyeball inspections on objects in the catalogue. The catalog includes Petrosian magnitudes and structural parameters in *ugriz* bands such as the concentrations $C = R_{90}/R_{50}$, where R_{90} and R_{50} are the Petrosian radii enclosing 90% and 50% of total luminosity. Effective radius, $R_{\rm eff}$, and the Sérsic indices n are taken from B05, where, a 1D seeing-convolved Sérsic profile is fitted to galaxies. The adopted redshift range sets the completeness level at a Petrosian magnitude $r \simeq 18$ and at a surface brightness within R_{50} $\mu_{50} \simeq 24.5\,{\rm mag\,arcsec^{-2}}$. This luminosity limited sample still suffers from an incompleteness at low surface brightness levels, e.g. in the regime of dwarf galaxies with $M_* \lesssim 10^{8.3-8.5}\,M_\odot$ (Blanton et al. 2005a, Baldry, Glazebrook & Driver 2008, Li & White 2009). Moreover, the incompleteness, almost negligible for more massive galaxies, comes back again for the most massive ones ($M_* \gtrsim 10^{11}\,M_\odot$) as the sample is biased against large objects at lower redshifts in consequence of a bias in the SDSS photometric pipeline (Blanton et al. 2005a). We address the systematics of the SDSS sample more in detail in App. A1.

Petrosian magnitudes are reduced to $z = 0$ (K-term; Blanton et al. 2003a), corrected for galactic extinction using the Schlegel, Finkbeiner & Davis (1998) dust map, converted to our adopted cosmology (Blanton et al. 2003b), and further corrected for the missing flux of Petrosian magnitudes. The latter correction is almost unchanged using the relation in Graham et al. (2005).

The sample spans a wide range of luminosities, masses, and colors, including galaxies lying both in the red-sequence (RS) and in the blue cloud (BC). In particular, we have sorted out ETGs by adopting the following two criteria:

- the Sérsic index satisfies the condition $2.5 \leqslant n \leqslant 5.5$;
- the concentration index $C > 2.6$ (Shimasaku et al. 2001, Padmanabhan et al. 2004, Cardone et al. 2008)[3].

[2] The catalog is available at: `http://sdss.physics.nyu.edu/vagc/lowz.html`.

[3] In order to retain the low mass red systems, no lower cut has

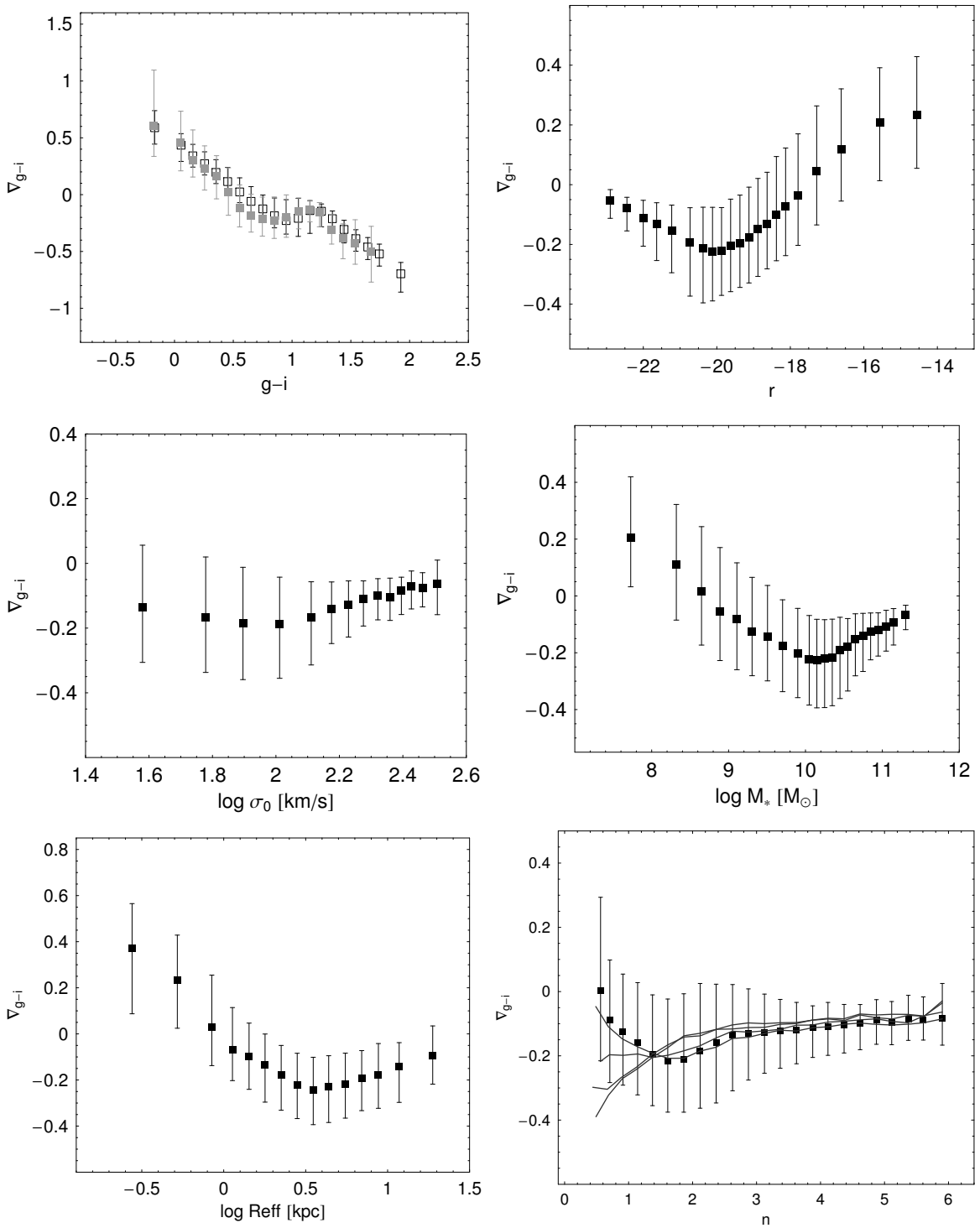

Figure 1. Gradients of the $g-i$ colour, as a function of various observed and derived quantities. Medians of binned CGs are plotted together with the $25-75\%$ quantiles, shown as error bars. *Top-left.* CG as a function of color, black open and gray squares are for colors at R_1 and R_2, respectively. *Top-right.* CG as a function of r-band magnitude. *Middle-left.* CG as a function of logarithm of σ_0, which is defined as the velocity dispersion within a circular aperture of radius $0.1R_{\rm eff}$ (using r-band $R_{\rm eff}$) from the SDSS velocity dispersion σ_{ap}, using the relation in Jørgensen et al. (1995, 1996). *Middle-right.* Gradient as a function of total stellar mass M_* (assuming a Chabrier IMF), which is the output of our stellar population analysis fitting the total colors. *Bottom-left.* CG as a function of logarithmic of r-band $R_{\rm eff}$ (the results does not depend on the band). *Bottom-right.* CG as a function of i-band Sérsic index n; the lines are the the median trends for the Sérsic index in the other bands.

The final ETG sample consists of 10 508 galaxies. Of the remaining entries of our database, 27 813 are LTGs, defined as objects with $C \leqslant 2.6$, $n \leqslant 2.5$, and $\sigma_0 \leqslant 150$ km/s, and 9 359 are intermediate morphology galaxies, with $C \leqslant 2.6$ (> 2.6) and $n > 2.5$ ($\leqslant 2.5$), hereafter named Intermediate-type galaxies (ITGs)[4]. Further details on the adopted morphology criteria are given in the App. A2. In this paper we will consider ETGs and LTGs only, with a main focus on ETGs.

2.2 Spectral models and gradients

The stellar population analysis of the selected sample is based on a set of "single burst" synthetic stellar spectra using the prescription of Bruzual & Charlot (2003, hereafter BC03). These models are widely used in literature both for ellipticals and spirals (e.g., Trager et al. 2000, Bell & de Jong 2001, Ganda et al. 2007, Peletier et al. 2007), but we are aware that, in some cases, e.g. in later-type systems with very protracted star formations (Kennicutt et al. 1983), they may fail to reproduce correctly the galaxy properties (MacArthur et al. 2009). A finer analysis would require a more realistic SF history for LTGs (e.g., an exponentially decaying or a constant SF), which is beyond the scope of this paper.

We adopt a Chabrier (2001, 2002, 2003) initial mass function (IMF), noting that the choice of a Salpeter (1955) IMF would affect the total stellar mass only, shifting $\log M_*$ upward by ~ 0.25 and leaving all the other stellar population parameters unchanged (e.g. Tortora et al. 2009a). Our choice of the BC03 models is driven by their high versatility and ability to span the stellar parameter space, e.g. metallicities and ages. There are other options (Rettura et al. 2006, Kannappan & Gawiser 2007, Tortora et al. 2009a for a comparison of the results from different prescriptions), but none is found to be bias free (Maraston 2005, van der Wel et al. 2006, Maraston et al. 2009, Conroy et al. 2009, 2010a,b).

In our model procedures we adopt a metallicity ranging within $0.0001 - 0.05$, and the Universe age, $t_{\rm univ}$ (in the assumed cosmology), as an upper limit for the galaxy ages. Synthetic colors have been obtained after convolving the spectra with the corresponding SDSS filter bandpass functions. To improve the sensitivity of the results to the small differences in both Z and ages, we have adopted a "mesh refinement" procedure which interpolates the synthetic models over the initial grid of equally spaced bins in the logarithm of Z and age. To estimate the galaxy age, Z, and the stellar mass-to-light ratio, Υ_*, we use a χ^2 minimization fit (see Tortora et al. 2009a,b for details) to the observed colors ($u-g$, $g-r$, $g-i$, and $g-z$). It is known that the use of just the optical broad bands is severely prone to the well known age-metallicity degeneracy. However, following previous indications (e.g. Wu et al. 2005) we have verified in App. B that the use of near infrared (NIR) photometry reduces the overall uncertainties but does not modify the estimates of the intrinsic population parameters based on optical bands only. We shall return on this issue later on.

We have used the structural parameters given by B05 to derive the color profile $(X-Y)(R)$ of each galaxy as the differences between the (logarithmic) surface brightness measurements in the two bands, X and Y. The CG is defined as the angular coefficient of the relation $X-Y$ vs $\log R/R_{\rm eff}$, $\nabla_{X-Y} = \dfrac{\delta(X-Y)}{\delta \log(R/R_{\rm eff})}$, measured in mag/dex (omitted in the following unless needed for clarity). The fit of each color profile is performed in the range $R_1 = 0.1R_{\rm eff} \leqslant R \leqslant R_2 = R_{\rm eff}$. We have verified that a choice of different radial ranges for the fit leaves the global trends unaffected. In particular, by setting $R_2 = 2R_{\rm eff}$ the gradients changed by $\lesssim 0.01 - 0.02$ mag, where $R_{\rm eff}$ is the r-band effective radius (Peletier et al. 1990b, La Barbera et al. 2005). Slightly different definitions are also used elsewhere (see discussion in Liu et al. 2009).

By definition, a positive CG, $\nabla_{X-Y} > 0$, means that a galaxy is redder as R increases, while it is bluer outward for a negative gradient. The fit of synthetic colors is performed on the colors at R_1 and R_2 and on the total integrated colors. Following La Barbera et al. (2005) we define the stellar parameters gradients as $\nabla_{\rm age} = \log[{\rm age}_2/{\rm age}_1]$ and $\nabla_{\rm Z} = \log[{\rm Z}_2/{\rm Z}_1]$, where (age_i, Z_i) with $i = 1, 2$ are the estimated age and metallicity at R_1 and R_2, respectively.

A full test of the reliability of our modelling technique, the presence of spuriously generated correlations, and the contribution from internal dust extinction are addressed in App. B.

3 RESULTS

We show how color and stellar population gradients depend on structural parameters and mass, and compare the results with the predictions from galaxy formation scenarios. The median trends, e.g. in bins of mass, velocity dispersion, colors etc., will be represented along with the 25-75th percentiles in each bin.

The average behaviour of galaxies is analyzed in the parameter space, and the galaxy properties discussed after splitting the sample in the two morphological classes, ETGs and LTGs. We concentrate mainly on the analysis of ETGs and compare our findings with a wide collection of independent data and simulations.

3.1 CGs as a function of structural parameters, luminosity and mass

Fig. 1 presents the gradients $\nabla_{\rm g-i}$ as a function of integrated colors, structural parameters, r-band magnitude, central velocity dispersion, and stellar mass. We adopted the $g-i$ color as a reference having varified that the use of different colors does not affect our conclusions. As shown in the top-left panel of the figure, CGs are, on average, positive for the galaxies with bluer colors, $g-i \lesssim 0.5$, while they reach increasingly negative values for redder galaxies (e.g. Park & Choi 2005; Lee et al. 2007; Suh et al. 2010). Since $g-i \sim 0.5$ marks the transition from the RS to the BC, this result allows us to broadly associate negative and positive

been applied to the velocity dispersion. However, very low values ($\sigma_0 < 70$ km/s) may be prone to large systematics, due to the signal-to-noise ratio and instrumental resolution of the SDSS spectra.

[4] Note that 2120 galaxies are cut away by the conservative upper limit on n, dictated by the fitting Sérsic code limit of $n = 6$.

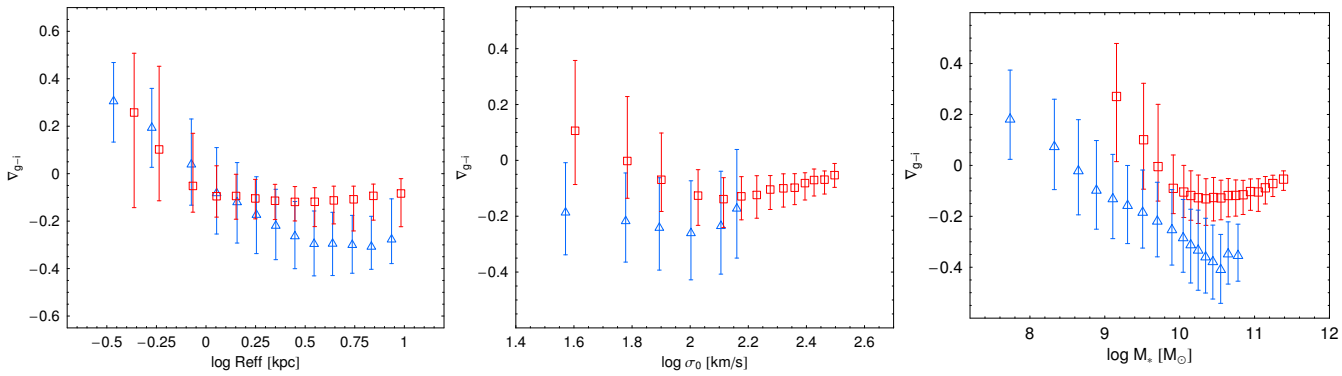

Figure 2. ∇_{g-i} as a function of effective radius (left panel), velocity dispersion (middle panel) and stellar mass (right panel) for ETGs (red symbols) and LTGs (blue symbols).

Table 1. Slopes of the correlation between CGs (∇_{u-g}, ∇_{g-r}, ∇_{g-i}, ∇_{g-z}), ∇_{age} and ∇_Z vs $\log\sigma_0$ and $\log M_*$ for ETGs and LTGs. The errors on slopes are computed as 1σ uncertainty by a bootstrap method. * For the LTG sample, splitting galaxies by masses smaller/larger than $\log M_* \sim 10.3$ would result in slopes of -0.07 ± 0.01 and 0.26 ± 0.25, respectively.

Correlation	γ_{ETG}		γ_{LTG}
	$\log\sigma_0 < 2.2$	$\log\sigma_0 > 2.2$	
$\nabla_{u-g} - \log\sigma_0$	-0.60 ± 0.03	0.40 ± 0.05	0.02 ± 0.05
$\nabla_{g-r} - \log\sigma_0$	-0.27 ± 0.01	0.19 ± 0.01	-0.02 ± 0.02
$\nabla_{g-i} - \log\sigma_0$	-0.34 ± 0.02	0.25 ± 0.04	-0.03 ± 0.02
$\nabla_{g-z} - \log\sigma_0$	-0.40 ± 0.02	0.17 ± 0.04	-0.09 ± 0.03
$\nabla_{age} - \log\sigma_0$	-0.10 ± 0.03	0.04 ± 0.17	0.14 ± 0.02
$\nabla_Z - \log\sigma_0$	-0.86 ± 0.07	1.02 ± 0.24	0.06 ± 0.10
	$\log M_* < 10.3$	$\log M_* > 10.3$	
$\nabla_{u-g} - \log M_*$	-0.43 ± 0.04	0.12 ± 0.02	-0.30 ± 0.01
$\nabla_{g-r} - \log M_*$	-0.24 ± 0.02	0.05 ± 0.01	-0.13 ± 0.01
$\nabla_{g-i} - \log M_*$	-0.34 ± 0.02	0.07 ± 0.01	-0.18 ± 0.01
$\nabla_{g-z} - \log M_*$	-0.36 ± 0.03	0.02 ± 0.01	-0.18 ± 0.01
$\nabla_{age} - \log M_*{}^\star$	-0.18 ± 0.06	0.24 ± 0.03	-0.01 ± 0.01
$\nabla_Z - \log M_*$	-0.46 ± 0.06	0.10 ± 0.04	-0.45 ± 0.01

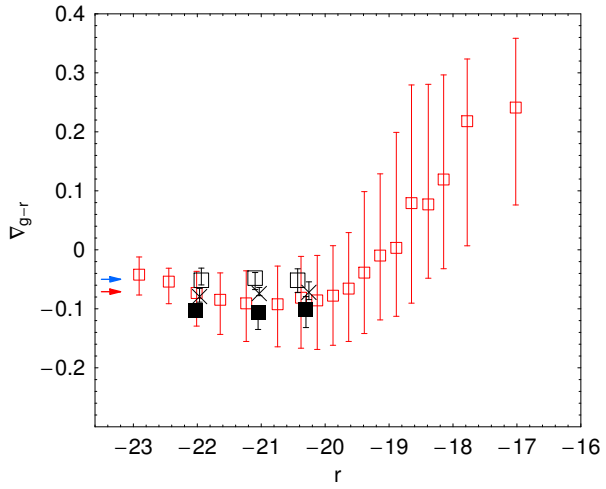

Figure 3. CG ∇_{g-r} as a function of r-band magnitude for ETG sample. Superimposed are the results from La Barbera et al. (2005) for high-richness (open boxes), low-richness (filled boxes), and all (crosses) the clusters. The blue and red arrows are the median ∇_{g-r} for the massive galaxies analyzed in Wu et al. (2005) and La Barbera et al. (2009) respectively; see the text for details.

CGs to the RS and BC, respectively. Furthermore, galaxies with intermediate colors ($g - i \sim 0.5 - 0.7$ or equivalently $g - r \sim 0.2 - 0.6$) have almost null (or slightly negative) gradients. The same trend holds by considering central and outer colors (at $R = R_1$ and R_2). We will return to this issue later on.

Tight trends are found when plotting the gradients against the r-band total magnitude and stellar mass, while the dependence on σ_0 is weaker. In the latter case, gradients look negative almost everywhere with the less massive galaxies ($\log\sigma_0 < 2\,$km/s) having the smallest ∇_{g-i} (~ -0.2). Although our sample does not include very massive galaxies ($M_* \gtrsim 10^{11.5}\,M_\odot$), our results seem to show a stable trend at the large mass scales, pointing to even shallower gradients. We note here that the smaller range shown by ∇_{g-i} as a function of σ_0 in the central panel (left column) of Fig. 1 is mainly due to the mix of the ETGs and LTGs. They have different trends with small σ_0 in the low mass regime (see also §3.2), thus producing a dilution of the average gradients with respect, e.g., the larger median ∇_{g-i} ($\sim 0 \div 0.2$) found for the low stellar mass (right column).

The correlations with luminosity and mass show that, starting from the bright/massive end, the CGs become steeper, with the steepest negative gradients (~ -0.2) corresponding to $r \sim -20$ mag and $M_* \sim 10^{10.3}\,M_\odot$. From there, bluer and later-type systems begin to dominate, and the gradients invert the trend with luminosity and mass, becoming positive at $r \sim -17.5$ mag and $M_* \sim 10^{8.7}\,M_\odot$. The lowest mass systems are those with the highest, positive gradients (~ 0.2). The mass scale at which the CG trend inverts is in remarkable agreement with the typical mass scale break for "bright" and "ordinary" galaxies (Capaccioli et al. 1992, see also §3.2) or for star-forming and passive systems (Kauffmann et al. 2003, Drory & Alvarez 2008).

The above trend holds also for the correlation with R_{eff} and, less strongly, with the Sérsic index, n, as expected in view of the tight correlation with stellar mass and luminosity (Caon, Capaccioli & D'Onofrio 1993, Prugniel & Simien 1997, Graham & Guzmán 2003, Shen et al. 2003, Trujillo et al. 2004, Tortora et al. 2009a). Note that the minimum at $n \sim 2$ for the trend of the gradient (Fig. 1, bottom right) roughly coincides with the value

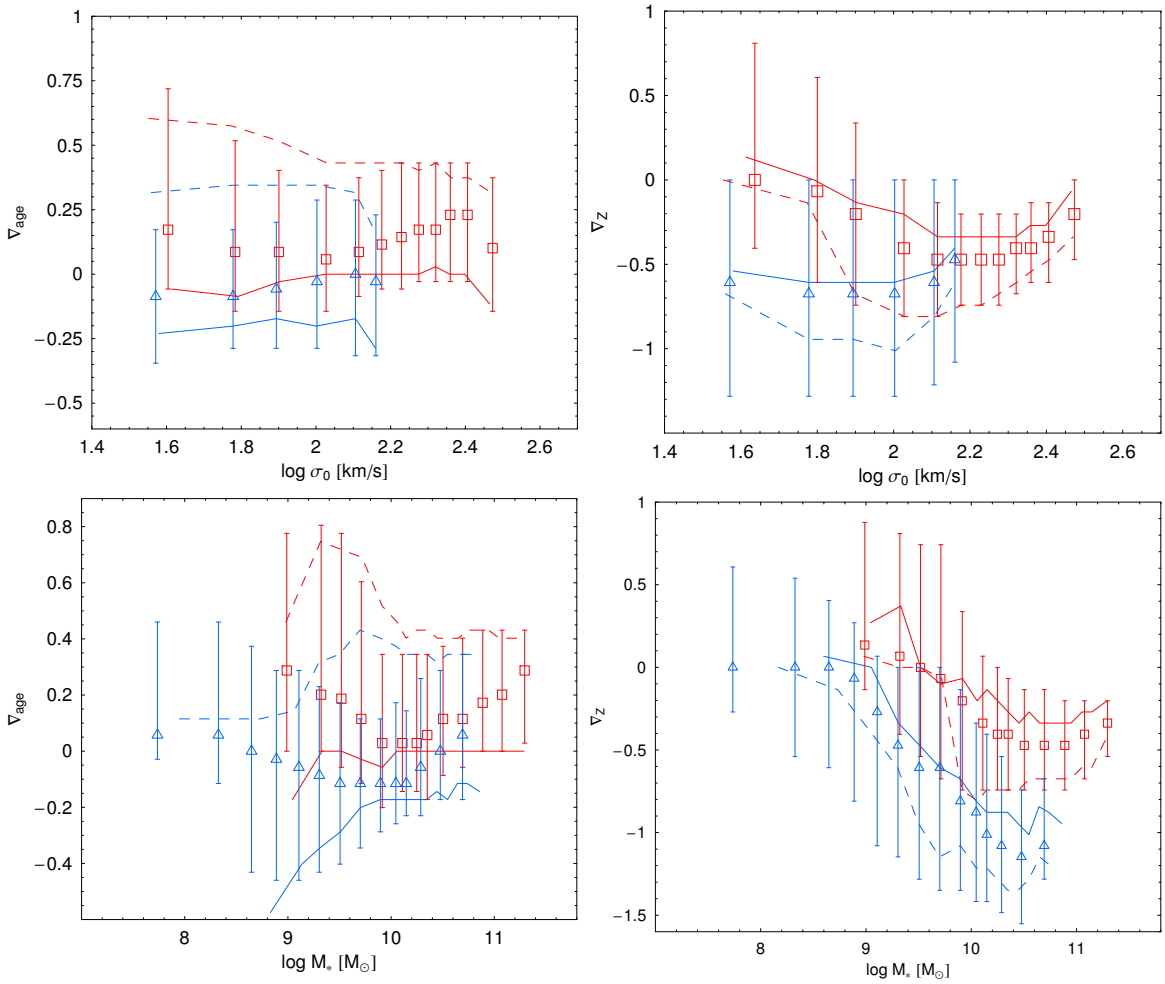

Figure 4. Age and metallicity gradients as a function of velocity dispersion (top panels) and stellar mass (bottom panels). Solid and dashed lines are for galaxies with central stellar populations older and younger than 6 Gyr, respectively. Symbols are as in Fig. 2

used here to set the separation between ETG and LTG systems, suggesting differences in the CG properties.

3.2 Gradients in morphologically selected galaxies

Following the dependence of ∇_{g-i} vs n just discussed, ETG and LTG samples will be now considered separately. In Fig. 2 we plot ∇_{g-i} versus $R_{\rm eff}$, σ_0, and M_* for ETGs (red symbols) and LTGs (blue symbols). The slopes of the linear best-fit relations of the CGs versus σ_0 and M_* are shown in Table 1.

First, ETGs are on average brighter and more massive than LTGs (Kauffmann et al. 2003) and have less dispersed CGs. As a general result, large $R_{\rm eff}$ galaxies are found to have lower gradients than more compact systems (e.g. Hopkins et al. 2009a). However, for LTGs, ∇_{g-i} monotonically decreases with $R_{\rm eff}$ and becomes flat at $\log R_{\rm eff} \gtrsim 0.5$, while for ETG it slightly increases for $\log R_{\rm eff} \gtrsim 0.3$, and decreases for smaller $R_{\rm eff}$. Similar trends are observed for the gradients as a function of the stellar mass. LTGs show a monotonic decreasing trend (in qualitative agreement with Kim & Ann 1990 and Liu et al. 2009), while a U-shaped function is found for ETGs with the gradients definitely decreasing with the mass for $M_* \lesssim 10^{10.3-10.5}\ M_\odot$, and a mildly increasing for larger mass values. This mass scale roughly corresponds to a total luminosity of $r \sim -20.5$ mag, which is compatible with the typical luminosity (and mass) scale for ETG dichotomy in the galaxy structural properties (e.g., Capaccioli et al. 1992a, Capaccioli, Caon & D'Onofrio 1992b, Graham et al. 2003). As discussed in Capaccioli et al. (1992a) this dichotomy may be related to the formation mechanism processes, and in particular to the dominance of the merging events for the "bright" sample, which is compatible with the shallower gradients that we observe for these systems (see, e.g., Fig. 2, right panel). Nevertheless, for a fixed stellar mass we observe that ETGs gradients are, on average, larger than LTG ones.

The same two-fold trend is shown for ETGs gradients as function of the velocity dispersion, while no trend with the mass is observed for LTGs gradients.

As a consistency check, in Fig. 3 we superpose to our $g-r$ gradients as a function of the $r-$band luminosity the results of La Barbera et al. (2005) obtained by analyzing a sample of low redshift luminous ($r \leqslant -20$ mag) cluster galaxies. Their value ($\nabla_{g-r} \sim -0.075$) is in agreement with our findings over the range where the two studies overlap, but we were also able to identify a change of ∇_{g-r} with our larger luminosity baseline. The results from the massive

galaxies in Wu et al. (2005) and La Barbera et al. (2009) are also shown to have ∇_{g-r} consistent with our results.

Unfortunately, we have no information on the environment of the individual systems in our sample, to check the dependence of our results with the local density. There are many lines of evidences that ETGs in clusters (and in very dense environments in general) have shallower gradients than systems in less dense environments (Tamura & Ohta 2000, Tamura et al. 2000, La Barbera et al. 2005 – see also Fig. 3). However, according to the morphological segregation, LTGs, which are expected to reside in the field or in the external regions of clusters, have indeed lower gradients than ETGs which are found to live mainly in higher density environments.

3.3 Interpretation in terms of Z and age gradients

Here we test whether the CGs can be explained in terms of stellar properties such as age and metallicity. As seen in §2, we have obtained the gradients of metallicity and age distributions using two characteristic scales ($0.1R_{\rm eff}$ and $1R_{\rm eff}$). The $\nabla_{\rm age}$ and $\nabla_{\rm Z}$ are shown in Fig. 4 as a function of the velocity dispersion and the stellar mass, and the fitted slopes of such relations are quoted in Table 1.

Some features survive the large scatter. For LTGs the $\nabla_{\rm age}$ is about zero both with σ_0 and mass, strongly suggesting that CGs should not depend on age gradients of the galaxy stellar population[5]. Actually, $\nabla_{\rm Z}$ is instead strongly dependent on M_*, with the lowest metallicity gradients (~ -1) at the largest masses, while the dependence on σ_0 is very weak as the $\nabla_{\rm Z} \sim (-0.5, -0.6)$ for all σ_0.

ETG age gradients seem basically featureless when plotted against σ_0 where we find $\nabla_{\rm age} > 0$ with typical median values of $\nabla_{\rm age} \sim 0.2$. On the other hand, the metallicity gradients show some features with $\nabla_{\rm Z}$ decreasing for $\log\sigma_0 \lesssim 2.2\,$km/s, with a slope $= -0.86$, and increasing for larger velocity dispersion, with a slope of 1.02, reaching the shallowest values (~ -0.2) at $\log\sigma_0 \gtrsim 2.4\,$km/s and at the very low σ_0 (i.e. $\log\sigma_0 < 1.8\,$km/s) where $\nabla_{\rm Z} \sim 0$. This peculiar trend is mirrored by a similar dependence on the stellar mass[6]. This two-fold behaviour is also significant when we plot $\nabla_{\rm age}$ as a function of the stellar mass: $\nabla_{\rm age}$ decreases at $\log M_* \lesssim 10.3-10.5$, and increases in more massive systems.

Bright ETGs with $r \lesssim -20$ mag ($\log M_* \gtrsim 10$) have a median $\nabla_{\rm age} = 0.144^{+0.004}_{-0.003}$ (with scatter $^{+0.26}_{-0.17}$) and $\nabla_{\rm Z} = -0.472^{+0.005}_{-0.004}$ (with scatter $^{+0.34}_{-0.27}$), where the error bars represent the 1σ uncertainty and the scatters the 25-75 percentiles of the sample distribution. However, we have found

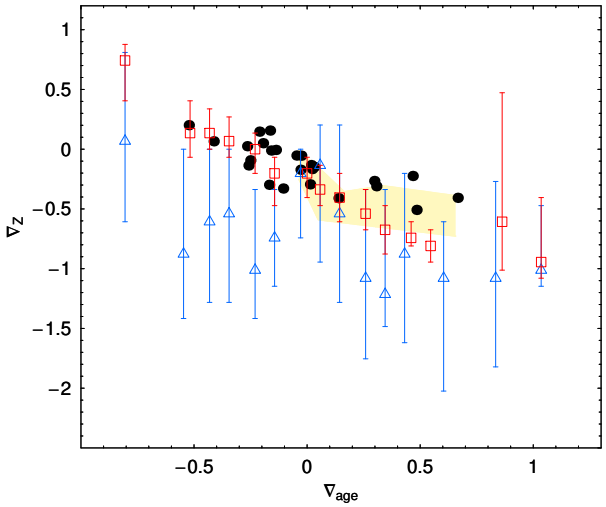

Figure 5. Metallicity vs age gradients for ETGs and LTGs. The symbols are the same as in Fig. 2 and 4. Metallicity vs age gradients for ETGs are compared with predictions from gas-rich mergings in Hopkins et al. (2009a) (yellow shaded region) and cluster ETGs in Rawle et al. (2009) (black points).

that $\nabla_{\rm age}$ and $\nabla_{\rm Z}$ strongly depend on galaxy age: in particular, if we separate the sample in systems with older and younger than 6 Gyr central stars, we obtain different age and metallicity gradient trends as shown in Fig. 4. Older systems have gradients which are shallower than younger (in particular for ETGs) and bracket the average trend of the whole samples. A partial contribution to the scatter is also given by the central metallicity: the lower-Z systems show generally shallower gradients and viceversa. However, the effect is significantly smaller with respect to the central age.

The median values of the older sample is $\nabla_{\rm age} = 0^{+0.003}_{-0.004}$ (with scatter $^{+0.14}_{-0.14}$) and $\nabla_{\rm Z} = -0.337^{+0.006}_{-0.005}$ (with scatter $^{+0.27}_{-0.20}$) while for the younger one we obtain $\nabla_{\rm age} = 0.43^{+0.001}_{-0.006}$ (with scatter $^{+0.08}_{-0.20}$) and $\nabla_{\rm Z} = -0.67^{+0.007}_{-0.004}$ (with scatter $^{+0.27}_{-0.13}$). Galaxies with old central ages are fully consistent with the results of pure metallicity and mixed age+metallicity models in La Barbera et al. (2005) and with the best fitted values in La Barbera et al. (2009) (where galaxy ages are in the range of $7.8 - 12.6\,$Gyr). The agreement with Rawle et al. (2009) is also good if we consider that their sample has a median central age of $\sim 10\,$Gyr (calculated within an $R_{\rm eff}/3$ aperture), comparable to the old objects of our sample: they find an age gradient of -0.02 ± 0.06 and metallicity gradients of -0.13 ± 0.04. Finally, from the analysis of 36 nearby ETGs with SDSS and 2MASS photometry, Wu et al. (2005) have derived $\nabla_{\rm age} = 0.02 \pm 0.04$ (and scatter 0.25) and $\nabla_{\rm Z} = -0.25 \pm 0.03$ (and scatter 0.19), which turn to be qualitatively consistent again with our "old" sample. The inclusion of dust extinction leaves our general considerations unaffected, as we will show in App. B3. The average trends of the metallicity and age gradients are almost unchanged, with variations going in the direction of a better match with other literature data.

The massive ETGs, with young cores ($< 6\,$Gyr) have very steep metallicity gradients, which are discrepant with respect to the literature results we have compared with. In App. A2 we have checked that this result does not depend

[5] The scatter of the gradients around zero has been suggested to be a manifestation of possible differences in the formation processes in these systems MacArthur et al. (2009).

[6] For ETGs, stellar mass and σ_0 are tightly correlated, thus the correlation and the interpretation discussed here can be made when dissecting the similar trend as a function of velocity dispersion. On the contrary, many low mass LTGs have very large velocity dispersion which contribute to have almost constant metallicity gradients for $\log\sigma_0 \lesssim 2.2\,$km/s and increasing at larger σ_0. But in this case the scatter is very large suggesting that the central velocity dispersion is not a representative parameter of such rotational velocity supported systems.

on the contamination of the ETG sample from late-type systems. These latter, indeed, also show the same dependence of gradients on central age (see Fig. 4). A clear example of systems with very steep gradients are the isolated galaxy NGC 821 which have a young central stellar population and a metallicity gradient of $\nabla_{\rm Z} = -0.72 \pm 0.04$ or NGC 2865, for which is $\nabla_{\rm Z} = -0.47 \pm 0.05$ (Proctor et al. 2005, Reda et al. 2007). One reason why we have been able to pick up this difference between centrally old and young systems is that most of the samples analyzed in literature are made up of cluster ellipticals (e.g. Mehlert et al. 2003, Rawle et al. 2009), with only few cases of field galaxies (Ogando et al. 2005, Reda et al. 2007), while our sample is composed by galaxies living in various environments. On average, field galaxies are found to be younger and have a wider distribution of galaxy age (Bernardi et al. 1998, Trager et al. 2000, Thomas et al. 2005), thus, a large fraction of our centrally young galaxies would be made of systems in the field or low density environments with steeper age and metallicity gradients than those in cluster galaxies (due to the little interactions with the neighborhood, see Sect. 3.2 and La Barbera et al. (2005)).

Fig. 5 displays a relevant anti-correlation between age and metallicity gradients: for ETGs the correlation is clear with a slope of -0.94 ± 0.02, while for LTGs it seems poorer but still significant, with a slope -0.48 ± 0.03 (see also Table 2). This is an intrinsic property of the galaxy sample, which is consistent with predictions from the galaxy merging scenario (e.g. Hopkins et al. 2009a) and with other observational evidences (e.g. Rawle et al. 2009). It is not spuriously induced by the well known age–metallicity degeneracy (see App. B). In particular, galaxies with strong negative metallicity gradients have positive age gradients (coherently with the merging simulation in the figure), while the occurrence of negative age gradients is associated to small mass systems with positive metallicity gradients. The negative correlation found for LTGs is also qualitatively consistent with the literature on spiral galaxies (MacArthur et al. 2004).

In Fig. 6 we finally show the correlations between gradients and stellar parameters at different galaxy radii, with their related slopes listed in Table 2. Metallicity and age gradients show a tight correlation with central metallicity and age, respectively. For the ETGs, the age gradients are generally different from zero, positive for small central ages, and null or negative for old systems. Metallicity gradients are negative for systems with older stellar populations at $R_{\rm eff}$ or with higher metallicity, while they are about zero for systems with old central ages or intermediate ages at $R_{\rm eff}$ ($4-7$ Gyr). Noticeably, the very positive $\nabla_{\rm Z}$ and negative $\nabla_{\rm age}$ correspond to very young populations at $R_{\rm eff}$, i.e. recent star formation outside the galaxy cores. These latter systems seem compatible with dwarf galaxies hosting expanding shells (e.g. Mori et al. 1997, see also §4).

Similarly, LTGs have also positive age gradients in presence of young cores, but in these cases we observe also large central metallicities (see Fig. 6, top row, third panel from left) and for these systems we also found the strongest metallicity gradients (bottom raw, third panel of Fig. 6).

Overall, results in Figs. 4 and 6 are in substantial agreement with the correlations reported in Rawle et al. (2009) [7] and with the findings in hydrodynamical simulations of both elliptical and disc galaxies (Hopkins et al. 2009a, Sánchez-Blázquez et al. 2009).

3.4 Comparison with literature data

We now proceed to a more detailed comparison of our findings with a set of literature works which make use of a more sophisticated analysis, although usually associated to smaller samples. In particular, we concentrate on the ETG sample, taking the results for LTGs with the necessary caution, due to the larger degree of uncertainties arising from, e.g., 1) the simple assumption on the star formation recipe, 2) the assumption of no dust gradients (see also App. B3), and 3) the prediction from theoretical simulations of the formation of such composite and complex systems which are still in their infancy (MacArthur et al. 2009, Martínez-Serrano et al. 2009, Sánchez-Blázquez et al. 2009).

Due to our minimal choice of the wavelength baseline and the possible biases on the stellar population estimates caused by the use of the optical band only, the comparison of our results with literature data is critical to 1) ultimately check the robustness of our results on the regions of the parameter space overlapping with independent analysis and 2) assess the gain in the physical information allowed by our analysis. In Fig. 7 we summarize the metallicity and age gradients for ETGs as a function of central velocity dispersion, and compare these with the observed gradients from a collection of literature studies (Mehlert et al. 2003, Proctor 2003, Ogando et al. 2005, Reda et al. 2007, Sánchez-Blázquez et al. 2007, Koleva et al. 2009a,b, Spolaor et al. 2009, Rawle et al. 2009). In the next section we will expand this comparison to the predictions from cosmological simulations. When considering massive ETGs, our $\nabla_{\rm Z}$ are systematically steeper than literature data (see top-left of Fig. 7), although showing on average a lower scatter than dwarf galaxies.

As seen in Fig. 6 and also illustrated in Fig. 7 (top-right panel), a significant contribution to the scatter of the metallicity gradients is linked to the age of galaxy cores, with older galaxies having, on average, shallower gradients than younger systems. When considering objects with $age_1 > 6$ Gyr, the agreement with the other studies (generally dealing with old systems; see §3.3) is remarkably good. The same conclusion is reached when considering the $\nabla_{\rm age}$

[7] In particular, if we limit our galaxies sample to luminous systems ($r < -20$ mag) with $Z_1 \geqslant 0.008$ and $age_1 \geqslant 2$ Gyr and perform the fit with respect to the central quantities within $R_{\rm eff}/3$ we find that the slope of the $\nabla_{\rm Z} - \log Z_1$ relation is -0.58 ± 0.04 in good agreement with the values in the range $[-0.6, -0.7]$ from Rawle et al. (2009). For $\nabla_{\rm age} - \log age_1$ relation we find a slope of -0.90 ± 0.06, which disagree with the shallower values reported in Rawle et al. (2009). A similar disagreement is found for the $\nabla_{\rm Z} - \log age_1$ and $\nabla_{\rm age} - \log Z_1$ correlations, where we find slopes 0.67 ± 0.05 and 0.50 ± 0.03 respectively, which are still not perfectly consistent with their results. These might be related to the small sample adopted, to the selected environment (i.e. cluster cores), to differences in population markers (i.e. line strengths rather than colors) and the use of synthetic spectral models with variable α-enhancement over-abundance.

Table 2. Slopes of the linear correlations between stellar population gradients ∇_Z, ∇_{age}, Z and *age* for ETGs and LTGs. IN and OUT indicate gradients correlated with inner (at $R = R_1$) and outer (at $R = R_2$) metallicity and age. In the last row the result for correlation between ∇_Z and ∇_{age} are reported. The fitting procedure is the same as that in Table 1. $^{(*)}$: here the linear correlation is a 0-th order approximation.

Correlation	γ_{ETG}		γ_{LTG}	
	IN	OUT	IN	OUT
$\nabla_{age} - \log age$	-0.80 ± 0.03	0.93 ± 0.03	-0.78 ± 0.02	0.75 ± 0.03
$\nabla_{age} - \log Z^{(*)}$	0.09 ± 0.01	-0.34 ± 0.01	0.21 ± 0.01	-0.24 ± 0.02
$\nabla_Z - \log age$	0.62 ± 0.05	-1.07 ± 0.04	-0.20 ± 0.06	-0.62 ± 0.04
$\nabla_Z - \log Z$	-0.64 ± 0.02	0.06 ± 0.02	-0.49 ± 0.01	0.66 ± 0.02
$\nabla_Z - \nabla_{age}$	-0.92 ± 0.02		-0.50 ± 0.03	

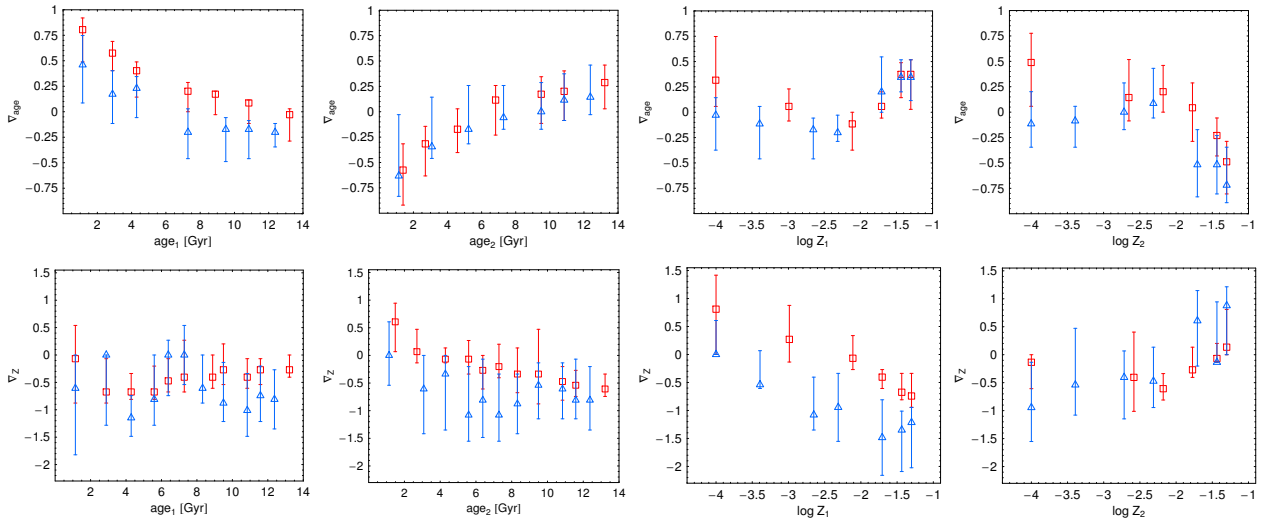

Figure 6. Correlations of gradients and fitted stellar parameters. The symbols are the same as in Figs. 2, 4 and 5. *Top panels.* From left, age gradient as a function of age_1, age_2, Z_1 and Z_2. *Bottom panels.* From left, metallicity gradient as a function of age_1, age_2, Z_1 and Z_2.

(Fig. 7, bottom-right), where the steeper (positive) gradients are shown by the younger systems. In App. B3 we show that shallower age gradients might be found if the dust extinction is accounted in the stellar models for the massive ETGs, while metallicity gradients remain almost unchanged.

The dependence of the gradients on the galaxy age is consistent with results in the dwarf regime by Spolaor et al. (2009) and Koleva et al. (2009a,b), which find shallower (and tightly correlated with mass) and steeper gradients in old and young systems, respectively.

As a final remark, galaxies with very low velocity dispersion ($\log \sigma_0 \lesssim 1.85$ km/s, see the vertical dashed gray line in Figs. 7 and 8) could be affected by some systematics due to the low signal-to-noise ratio and instrumental resolution of SDSS spectra. Notwithstanding this caution, our recovered gradients reproduce fairly well the literature results for dwarf ETGs (e.g. Spolaor et al. 2009). Also, the indication of shallower gradients at the lowest σ_0 seems robust, due to the clear turnover at $\log \sigma_0 \sim 2.2$ km/s and the inversion in the range $1.85 \lesssim \log \sigma_0 \lesssim 2.2$ km/s.

4 DISCUSSION

Our results seem to support the idea that the metallicity trend versus the stellar mass for LTGs is mainly driven by the interplay of gas inflow and winds from supernovae and evolved stars (Matteucci 1994, Kobayashi 2004, Pipino et al. 2008). These processes tend to increase the central metallicity and prevent the enrichment of the outer regions. Therefore more massive systems have on average larger central metallicities which correspond to steeper negative gradients.

Low mass ETGs ($M_* \lesssim 10^{10.3-10.5}\, M_\odot$, see Table 1) show a similar correlation of the metallicity gradient with stellar mass (Fig. 4), which suggests that they might experience the effect of infall/SN feedback as LTGs. For instance, low mass ETGs in Fig. 8 are consistent with SN feedback (both soft and strong), as predicted in chemo-dynamical simulations in Kawata (2001). On average, weaker SN feedback gives gradients in agreement with ours at intermediate $\log \sigma_0$ (~ 2.2 km/s), while the stronger feedback recipe seems to reproduce better the low $\log \sigma_0$ side of the correlation. A recipe including the SN feedback with a varying power as a function of mass (Dekel & Silk 1986, Dekel & Birnboim 2006) would allow a fairly good match of the observed decreasing trend of the gradients for the less massive galaxies

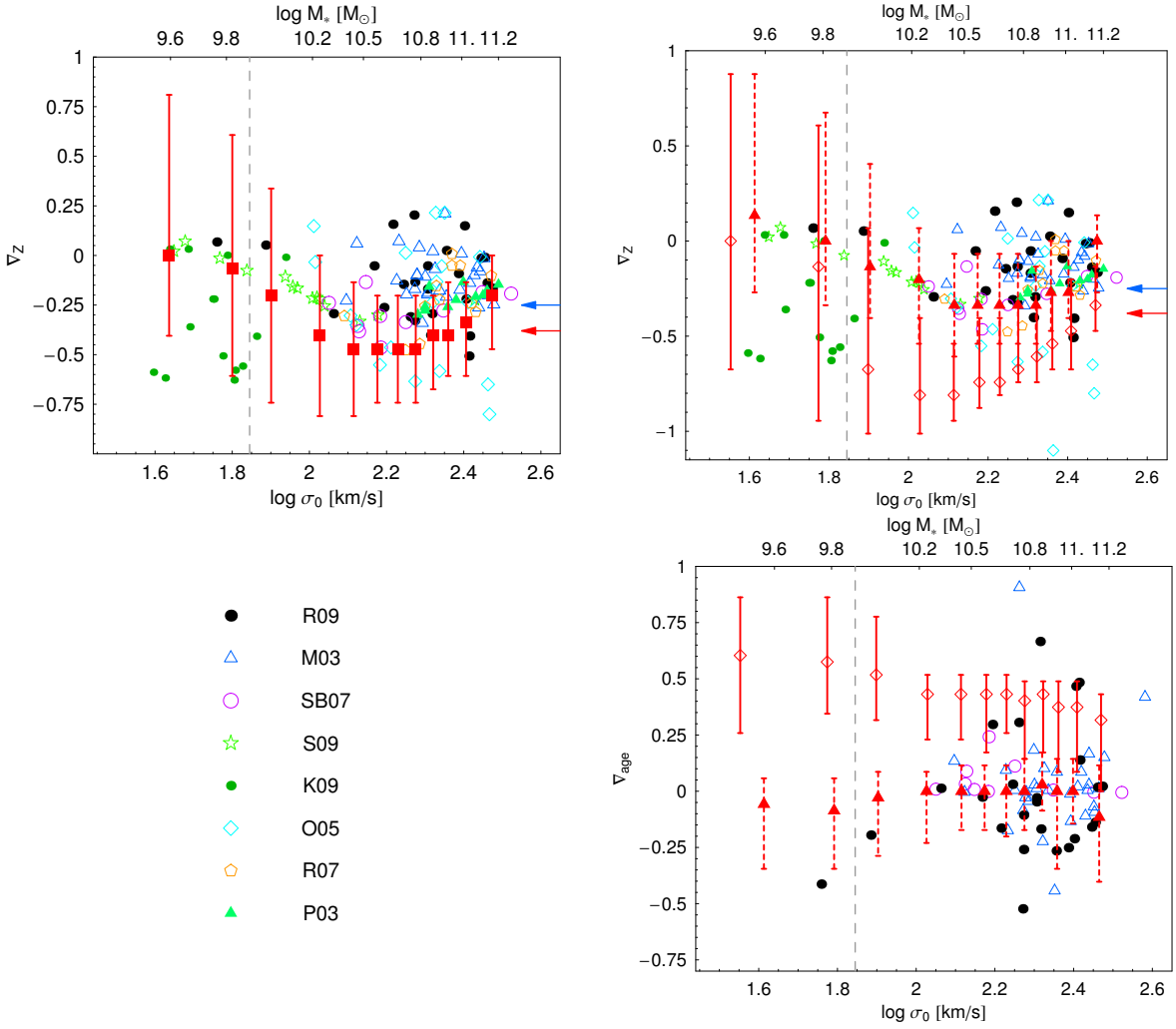

Figure 7. Metallicity gradients for ETGs as a function of central velocity dispersion compared with some literature data. Galaxies with velocity dispersion $\lesssim 70$ km/s, at the left of the vertical dashed gray line, are somehow uncertain. *Top-left.* Metallicity gradients are compared with Rawle et al. (2009) (R09) as black points (including the results for Shapley and A3389 clusters), Mehlert et al. (2003) (M03) as blue triangles, Sánchez-Blázquez et al. (2007) (SB07) as purple circles, Spolaor et al. (2009) (S09) as green stars, Koleva et al. (2009a,b) (K09) as light green points, Ogando et al. (2005) (O05) as cyan diamonds, Reda et al. (2007) (R07) as orange pentagons and Proctor (2003) (P03) as green triangles. The blue and red arrows are for the results in Wu et al. (2005) and La Barbera et al. (2009). *Top-right.* Our ETGs divided in young and old central ages are compared with literature data: diamonds with continue bars are for ETGs with $0 < age_1 \leqslant 6$ Gyr; triangle with dashed bars are for $age_1 > 6$ Gyr. *Bottom-left.* Legend of symbols for literature data. *Bottom-right.* Age gradients divided by age are plotted versus literature data (M03, SB07 and R09) as above.

(as in Kawata & Gibson 2003), but it fails to reproduce our observed gradients for the very massive systems. Moreover, as seen in Fig. 4, the absolute value of ∇_Z of ETGs is lower than for the LTGs, probably as a consequence of the dilution of the gradient due to the higher-density environment where ETGs generally live in (see also discussion in §3.2). It might be also the effect of other mechanisms at work such as merging (see, e.g., Kobayashi 2004), which however is expected to be less effective in these mass ranges (de Lucia et al. 2006).

At very low masses ($M_* \lesssim 10^{9.5}\,M_\odot$ and $\log\sigma_c \lesssim 1.8$ km/s) the ∇_Z turns to even positive values, which mainly corresponds to systems with negative ∇_{age} as in Fig. 5 and very low central metallicities as in Fig. 6: these dwarf-like systems are compatible with the expanding shell model from Mori et al. (1997) (see also Fig. 8).

At larger masses, the shallow metallicity gradients of ETGs suggest that these have experienced merging and/or tidal interactions (at a rate that could increase as a function of the stellar mass) which have diluted the ∇_Z. Such events might have taken place in earlier phases of the galaxy evolution, as indicated by the presence of null age gradients in old systems. In Fig. 8 we compare the ∇_{age} to the prediction from the hierarchical simulations in Hopkins et al. 2009a, where a perfect agreement is found for the old sample (see also the good agreement with literature data if Fig. 7). From the same simulations we also see that systems with younger cores are expected to show positive gradients which are fully consistent with our low central age systems as in Fig. 4 and also reported in Fig. 8. In the same figure, metallicity gradients are in qualitative agreement with gradients of remnants in Hopkins et al. (2009a) and the medians from Kobayashi (2004), while the merg-

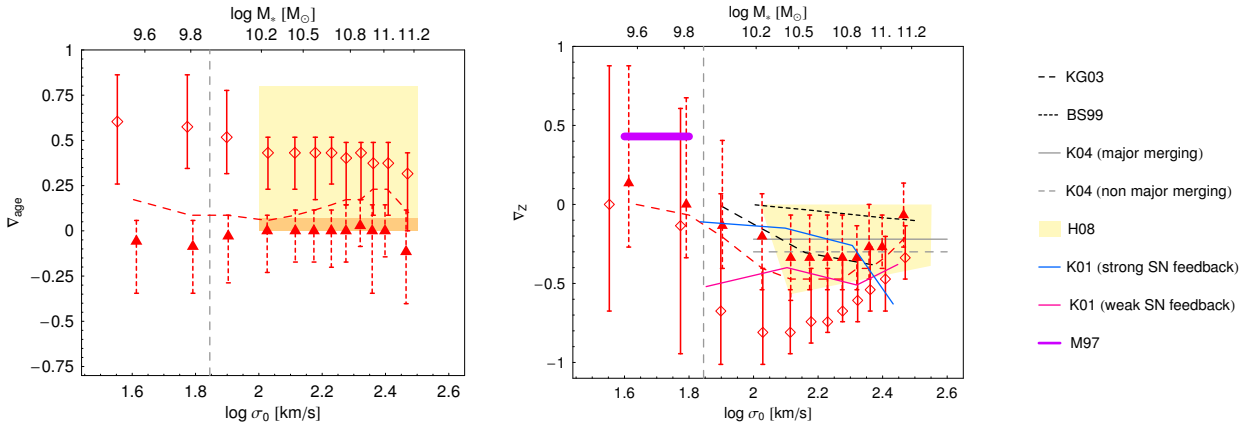

Figure 8. *Left.* Age gradients versus simulations. Comparison with gas-rich merging predictions in Hopkins et al. (2009a), with the darker (lighter) region showing galaxies older (younger) than 6 Gyr. *Middle.* Metallicity gradients versus simulations. The predictions from merging models in Bekki & Shioya (1999) are shown as short dashed line, dissipative collapse models in Kawata & Gibson (2003) as long dashed line, the remnants of major-mergings between gas-rich disk galaxies in Hopkins et al. (2009a) as yellow shaded region, the typical gradient for major (continue gray line) and non-major (dashed gray line) mergings in Kobayashi (2004). The violet thick line shows the predictions as in the simulation of Mori et al. (1997); blue and pink lines are the result from the chemo-dynamical model in Kawata (2001), respectively for strong and weak supernovae feedback models (the B-band magnitude in this work is transformed in a velocity dispersion σ_0 using the relation in Tortora et al. (2009a)). The symbols are as in Fig. 7 and the dashed red line is for the whole ETGs sample. *Right.* Symbol legend for simulations.

ing models in Bekki & Shioya (1999) reproduce the mean gradients in galaxies with $\log\sigma_0 \gtrsim 2.4$, but fail at lower σ_0. Once again, if we consider the older systems only, they are in a better agreement with the merging simulations: this is expected since, after the initial gas rich-merging events that might produce both a larger central metallicity and a positive age gradient (Kobayashi 2004, Mihos & Hernquist 1994), subsequent gas poor-merging may dilute the positive age gradient with time as well as make the metallicity gradients to flatten out (White 1980, Hopkins et al. 2009a, Di Matteo et al. 2009) – see also left column of Fig. 6. In this respect, massive and old ETG systems seem to be fully consistent with the merging scenario. However, a further mechanism that may act to produce the shallower (or almost null) color and metallicity gradients in old massive ETGs, with $M_* \gtrsim 10^{11}\,M_\odot$ and $\log\sigma_0 \gtrsim 2.4$ (Fig. 4), might be due to the strong quasar feedback at high redshift (Tortora et al. 2009b), while steeper metallicity gradients at lower masses could be linked to less efficient AGNs.

5 CONCLUSIONS

We have investigated the colour gradients in a sample of $\sim$ 50 000 local galaxies from the SDSS as a function of structural parameters, luminosity, and stellar mass. CGs have been found to correlate mainly with luminosity and stellar mass. They have a negative minimum ($\nabla_{g-i} \sim -0.2$) at $M_* \sim 10^{10.3}\,M_\odot$, and increase with the mass for $M_* \gtrsim 10^{10.3}\,M_\odot$, the very massive galaxies, $M_* \gtrsim 10^{11}\,M_\odot$, having the shallower values (~ -0.1). On the other mass side, the gradients decrease with mass and become positive for $M_* \lesssim 10^{9}\,M_\odot$. These trends are mirrored by similar behaviours with luminosity and galaxy size, e.g. the gradients have a minimum at $r \sim -20$ mag and $\log R_{\rm eff} \sim 0.5$, then increase toward the small and the large end of the parameter distribution, turning to positive values for $\log R_{\rm eff} \lesssim 0$ and $r \gtrsim -18$ mag. The dependence on velocity dispersion is very loose. A clear dichotomy is also found when looking at the dependence on the Sérsic index n which suggests a distinct behaviour between the ETG and LTG. In fact, these two families mark clear differences in their trends with structural parameters (e.g. mass and σ_0).

For LTGs, ∇_{g-i} monotonically decreases with the mass, with more massive systems having the lowest CGs (~ -0.4) and less massive systems ($M_* \lesssim 10^{8-8.5}\,M_\odot$) showing even positive gradients. ETGs have negative gradients mildly increasing with mass for $\log M_* \gtrsim 10^{10.3-10.5}\,M_\odot$, which marks the mass scale for the gradient slope inversion (see Table 1). This result is consistent with Spolaor et al. (2009) and reminiscent of the typical mass scale where the star formation and structural parameters in galaxies drastically change (Capaccioli et al. 1992a; Capaccioli, Caon & D'Onofrio 1992b, Kauffmann et al. 2003, Croton et al. 2006, Cattaneo et al. 2008). A similar trend is observed when plotting the gradients as a function of effective radius, while a tighter trend with velocity dispersion is evident for ETGs only (Fig. 4). This was masked when plotting LTGs and ETGs together.

We have used galaxy colors at $0.1R_{\rm eff}$ and $1R_{\rm eff}$ in our synthetic spectral models to determine the variation of age and metallicity at these radii. The observed trends of the CGs with mass and σ_0 are correlated with a similar trends for metallicity and age gradients. Despite the large scatter of the data, the strong correlation of metallicity gradients with the central velocity dispersion of ETGs is clear, with a turnoff point at $\log\sigma_0 \sim 2.2$ km/s (see Table 1). These results are in very good agreement with a collection of results from literature (Mehlert et al. 2003, Proctor 2003, Ogando et al. 2005, Reda et al. 2007, Sánchez-Blázquez et al. 2007, Koleva et al. 2009a,b, Spolaor et al. 2009, Rawle et al. 2009), in particular for the old galaxies in our sample.

A remarkable result of our analysis is the confirmation that the galaxy (central) age is one of the main drivers of the scatter of the age and metallicity gradients, with the older systems showing generally the shallower gradients with respect the young ones. In Fig. 7, for instance, we show that in the low mass regime, at fixed σ_0, older galaxies have on average shallower gradients, consistently with the results in Spolaor et al. (2009), while systems with late formation or with a recent SF episode have a larger spread (and metallicity gradients down to very low values, ~ -0.6), consistently with findings in Koleva et al. (2009a,b). The measured scatter might be the consequence of a variety of phenomena affecting the dwarf galaxy evolution, such as (soft) SN feedback, interaction/merging in the high density environment, and star-formation by shell expansion (Fig. 8). On the other mass side, a further factor of the spread of the massive systems gradients might be the randomness of the initial conditions of the mergings (Rawle et al. 2009).

Previous attempts to quantify the correlations of gradients with luminosity, mass, or velocity dispersion have often failed, mainly because of the exiguity of the galaxy sample (e.g. Peletier et al. 1990a, Kobayashi & Arimoto 1999, Tamura & Ohta 2003). Only recently a clear correlation of the color and metallicity gradients with mass has been ascertained (see e.g., Forbes et al. 2005), pointing to different trends for high and low mass galaxies (Spolaor et al. 2009, Rawle et al. 2009). Our analysis has confirmed and reinforced these results, as it relies on one of the largest local galaxy samples including dwarf, normal, and giant galaxies. Using CGs derived by the structural parameters in B05, we obtained statistically meaningful trends, which allow us to fix the link with physical phenomena as well as to make direct comparisons with predictions from simulations (Mori et al. 1997, Bekki & Shioya 1999, Kawata 2001, Kawata & Gibson 2003, Kobayashi 2004, Hopkins et al. 2009a).

We can finally draw the physical scenario sketched in Fig. 9, resting on the results discussed so far. The formation of LTGs and less massive ETGs is mainly driven by a monolithic-like collapse as they have almost null or positive age gradients and negative metallicity gradients decreasing with the mass. These are in fact well explained by simple gas inflow and feedback from SN and evolved stars, and also reproduced in simulations of dissipative collapse and SN feedback models in Fig. 8 (Larson 1974, 1975, Carlberg 1984, Arimoto & Yoshii 1987, Kawata & Gibson 2003). The difference in the magnitude of the gradients between LTGs and lower mass ETGs mainly resides in the effect of the environment as the former live in very low density environments, while the latter stay in massive haloes, together with more massive ETGs. Here they experience strong gravitational interactions producing shallower color and metallicity gradients. The efficiency of such phenomena is strong enough for lower mass ETGs, thus shaping the observed decreasing trend with mass (see e.g. Dekel & Birnboim 2006).

The most massive ETG systems have shallower gradients (Bekki & Shioya 1999, Kobayashi 2004, Hopkins et al. 2009a), flattening out with the mass due to the increasing intervention of (gas-rich and -poor) merging, tidal interactions, and quasar/radio mode AGN feedback (Capaccioli et al. 1992a, Ferrarese & Merritt 2000, de Lucia et al. 2006, Dekel & Birnboim 2006,

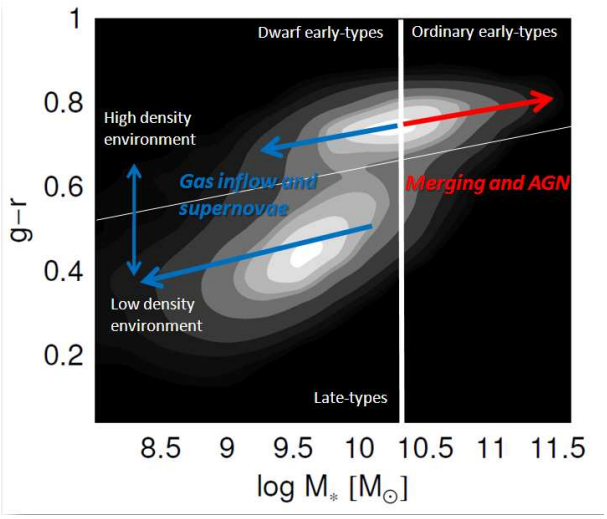

Figure 9. Color-mass diagram. The contours show the density of data-points with gray scale going from darker (low density of galaxies) to brighter regions (high density). The thin white line sets the separation between RS and BC; the thick vertical one gives the mass scale, $M_* \sim 10^{10.3}$, which separates galaxies belonging to the RS in normal and dwarf ETGs, and sets a qualitative upper mass for LTGs. The arrows give information about the efficiency of the phenomena which drive the two-fold trend we discuss in the paper. AGN feedback and merging are important at high mass with an efficiency increasing with mass, while SN feedback and gas inflow drive the galaxy evolution in the less massive side, with an efficiency that is larger at the lowest masses. Galaxies in RS and BC lie in environments with a different density, which is manifested as a difference of the gradients in the two samples.

Liu, Jiang, & Gu 2006, Sijacki et al. 2007, Cattaneo et al. 2008). All the simulations collected in literature fail to reproduce the fine details of the trends we find, suggesting that a full understanding of physical processes involved in the galaxy evolution is still missing.

In future analysis we plan to enlarge the wavelength baseline and also to use line-strength measurements for the stellar model (when available). We will investigate the systematics induced when other synthetic prescriptions are assumed and the dependence of the CGs on the galaxy star formation history. We will derive the M/L gradients, if any, and discuss their impact on the dark matter content of ETGs and correlate the CGs with the DM fraction of these systems as probe of the galaxy potential wells. Finally, the realization of hydrodynamical simulations of jets emitted by AGN would be useful to shape, together with galaxy merging, the color and metallicity gradients for massive galaxies.

ACKNOWLEDGMENTS

We thank the anonymous referee for the useful suggestions which helped to improve the paper and the robustness of the results. We also thank F. La Barbera for the fruitful discussion and comments. CT is supported by the Swiss National Science Foundation and by a grant from the project Mecenass, funded by the Compagnia di San Paolo. VFC is

supported by Regione Piemonte and Università di Torino and partially from INFN project PD51.

REFERENCES

Allen P. D., Driver S. P., Graham A. W., Cameron E., Liske J., De Propris R., 2006, MNRAS, 371, 2
Antonuccio-Delogu V. & Silk J. 2008, MNRAS, 389, 1750
Arimoto N. & Yoshii Y. 1987, A&A, 173, 23
Balcells M. & Peletier R. F. 1994, AJ, 107, 135
Baldry I. K., Glazebrook K., Driver S. P., 2008, MNRAS, 388, 945
Bakos J., Trujillo I. & Pohlen M. 2008, ApJ, 683, 103
Bekki, K., & Shioya, Y. 1999, ApJ, 513, 108
Bell Eric F. & de Jong R. S. 2001, ApJ, 550, 212
Bell, E. F., McIntosh, D. H., Katz, N., & Weinberg, M. D. 2003, ApJS, 149, 289
Bernardi M. et al. 1998, ApJ, 508, 143
Blanton M. R. et al., 2003a, AJ, 125, 2348
Blanton M. R. et al., 2003b, ApJ, 592, 819
Blanton M. R., Lupton R. H., Schlegel D. J., Strauss M. A., Brinkmann J., Fukugita M., Loveday J., 2005a, ApJ, 631, 208
Blanton M. R. et al. 2005, AJ, 129, 2562 (B05)
Brammer G. B. et al., 2009, ApJ 706, L173
Bruzual, A. G. & Charlot, S. 2003, MNRAS, 344, 1000
Bundy, K., Ellis, R.S., Conselice, C.J., Taylor, J.E., et al. 2006, ApJ, 651, 120
Burstein D., Faber S. M., Gaskell C. M., Krumm N., 1984, ApJ, 287, 586
Caon, N., Capaccioli, M. & D'Onofrio, M. 1993, MNRAS, 265, 1013
Capaccioli M., Caon N. & D'Onofrio M. 1992, MNRAS, 259, 323
Capaccioli M., Caon N. & DOnofrio M., 1992, ESO ESP/EIPCWorkshop on Structure, Dynamics and Chemical Evolution of Early-type Galaxies (Elba), eds. J. Danziger, W. W. Zeilinger, and K. Kjar, ESO: Garching, 43
Cappellari, M. et al. 2006, MNRAS, 366, 1126
Cardiel N., Gorgas J., Sánchez-Blázquez P., Cenarro A. J., Pedraz S., Bruzual G. & Klement J. 2003, A&A, 409, 511
Cardelli J.A., Clayton G.C., Mathis J.S. 1989, ApJ, 345, 245
Cardone V.F. et al. 2008, astro-ph/0703532
Carlberg R. G., 1984, ApJ, 286, 403
Carollo C. M., Danziger I. J. & Buson L., 1993, MNRAS, 265, 553
Cattaneo, A., Dekel, A., Devriendt, J., Guiderdoni, B., & Blaizot, J., 2006, MNRAS, 370, 1651
Cattaneo, A., Dekel, A., Faber, S. M., & Guiderdoni, B. 2008, MNRAS, 389, 567
Cattaneo, A., Mamon, G. A., Warnick, K., Knebe, A. 2010, arXiv1002.3257
Chaboyer, B., 1994, in ESOyOHP Workshop on Dwarf Galaxies., ed. G. Meylan & P. Prugniel (Garching: ESO), 485
Chabrier, G. 2001, ApJ, 554, 1274
Chabrier, G. 2002, ApJ, 567, 304
Chabrier, G. 2003, PASP, 115, 763
Chang et al. 2006, MNRAS, 366, 717
Chiosi C. & Carraro G. 2002, MNRAS, 335, 335
Cimatti A., Daddi E., Renzini A., 2006, A&A 453, L29
Conroy, C. & Wechsler, R. H. 2008, arXiv0805.3346
Conroy, C., Gunn, J. E., & White, M. 2009, ApJ, 699, 486
Conroy, C., Gunn, J. E., & White, M. 2010a, ApJ, 708, 58
Conroy, C. & Gunn J. E. 2009, arXiv:0911.3151
Cowie L. L., Songaila A., Hu E. M., Cohen J. G. 1996, AJ, 112, 839
Croton D. J. et al., 2006, MNRAS, 365, 11
Daddi E. et al. 2005, ApJ, 626, 680
Dalcanton J. J. & Bernstein R. A. 2002, AJ, 124, 1328
Davies R. L., Sadler E. M., Peletier R. F. 1993, MNRAS, 262, 650
de Jong R. S. 1996, A&A, 313, 377
de Lucia, G., Springel, V., White, S. D. M., Croton, D., Kauffmann, G. 2006, MNRAS, 366, 499D
Dekel A. & Silk J., 1986, ApJ, 303, 39
Dekel A., & Birnboim Y. 2006, MNRAS, 368, 2
Dorman B, O'Connell R.W. & Rood R. T. 2003, ApJ, 591, 878
Drory N. & Alvarez M. 2008, ApJ, 680, 41
di Matteo T., Springel V. & Hernquist L. 2005, Nature, 433, 604
di Matteo P., Pipino A., Lehnert M. D., Combes F., Semelin B. 2009, A&A, 499, 427
Ferrarese, L. & Merritt, D. 2000, ApJ, 539, 9
Fontanot, F. et al. 2009, arXiv:0901.1130
Forbes D.A. Sánchez-Blázquez P. & Proctor R. 2005, MNRAS, 361, 6
Franx, M., Illingworth, B., & Heckman, T. 1989, AJ, 98, 538
Gallazzi, A. et al. 2005, MNRAS, 362, 41
Ganda K. et al. 2007, MNRAS, 380, 506
Gibson B. K. 1997, MNRAS, 290, 471
Glazebrook K. et al. 2004, Nature, 430, 181
Graham A. W. & Guzmán R. 2003, AJ, 125, 2936
Graham, A. W., Erwin, P., Trujillo, I., & Asensio Ramos, A. 2003, AJ, 125, 2951
Graham, A. W. et al. 2005, AJ, 130, 1535
Goudfrooij P. & de Jong T. 1995, A&A, 298, 784
Guo Y. et al. 2009, MNRAS, 398, 1129
Henry R. B. C., Worthey G. 1999, PASP, 111, 919
Hopkins P. F., Cox T. J., Dutta S. N., Hernquist L., Kormendy J., & Lauer T. R. 2009a, ApJS, 181, 135
Hopkins P. F., Lauer T. R., Cox T. J., Hernquist L., & Kormendy J. 2009b, ApJS, 181, 486
Jansen R. A., Franx M., Fabricant D. & Caldwell N. 2000, ApJS, 126, 271
Jarrett T. H., Chester T., Cutri R., Schneider S. E., Huchra J. P. 2003, AJ, 125, 525
Jimenez R., Panter B., Heavens A. F., Verde L., 2005, MNRAS, 356, 495
Jørgensen, I., Franx, M., Kjærgaard, P. 1995, MNRAS, 273, 1097
Jørgensen, I., Franx, M., Kjærgaard, P. 1996, MNRAS, 280, 167
Kannappan, S. J. & Gawiser, E. 2007, ApJ, 657, 5
Kauffmann, G. & White, S. D. M., 1993, MNRAS 261, 921
Kauffmann G. et al. 2003, MNRAS, 341, 54
Kaviraj S. et al. 2007, MNRAS 382, 960
Kawata D. 2001, ApJ, 558, 598
Kawata D. & Gibson B. K. 2003, MNRAS, 340, 908

Kennicutt R. C., 1983, ApJ, 272, 54
Khalatyan A., Cattaneo A., Schramm M., Gottlber S., Steinmetz M., Wisotzki L, 2008, MNRAS, 387, 13
Kim K. O. & Ann, H. B. 1990, The Journal of the Korean Astronomical Society, 22, 43
Kormendy, J., & Djorgovski, S. 1989, ARA&A, 27, 235
Kobayashi C. & Arimoto N. 1999, ApJ, 527, 573
Kobayashi C. 2004, MNRAS, 347, 740
Koleva M. et al. 2009, MNRAS 396, 2133
Koleva M. et al. 2009, Astron. Nachr. / AN 330, No. 00, 960 (arXiv:0910.2643)
Kuntschner H., Emsellem E., Bacon R. et al. 2006, MNRAS, 369, 497
La Barbera, F., Busarello, G., Merluzzi, P., Massarotti, M., & Capaccioli, M. 2002, ApJ, 571, 790
La Barbera F., Busarello, G., Massarotti, M., Merluzzi, P., & Mercurio, A. 2003, A&A, 409, 21
La Barbera F., et al. 2005, ApJ, 626, 19
La Barbera F., et al. 2009, ApJ, 699, 76
Larson R. B., 1974, MNRAS, 166, 585
Larson R. B., 1975, MNRAS, 173, 671
Lee J. H., et al. 2007, ApJ, 663, 69
Li Z. Han Z. & Zhang F. 2007, A&A, 464, 853
Li C. & White S. D. M. 2009, MNRAS, 398, 2177
Lintott C. et al., 2008, MNRAS, 389, 1179
Liu Y., Jiang D. R. & Gu M. F. 2006, ApJ, 637, 669
Liu C. Z., et al. 2009, arXiv:0905.4020
MacArthur L. A., Courteau S., Bell E. & Holtzman J. A. 2004, ApJS, 152, 175
MacArthur L. A. 2005, ApJ, 623, 795
MacArthur L. A., González J. J. & Courteau S. 2009, MNRAS, 395, 28
Maller A. H., Berlind A. A., Blanton M. R., Hogg D. W., 2009, ApJ, 691, 394
Maraston C. 2005, MNRAS, 362, 799
Maraston C. et al. 2009, MNRAS, 394, 107
Martínez-Serrano F. J. , Serna A., Doménech-Moral M. & Domínguez-Tenreiro R. 2009, ApJ, 705, 133
Matteucci F. 1994, A&A, 288, 57
Mehlert D., Thomas D., Saglia R. P., Bender R., Wegner G., 2003, A&A, 407, 423
Mihos J. C. & Hernquist L., 1994, ApJ, 437, L47
Mori, M. 1997, ApJ, 478, 21
Nakamura O., Fukugita M., Yasuda N., Loveday J., Brinkmann J., Schneider D. P., Shimasaku K., SubbaRao M., 2003, ApJ, 125, 1682
Napolitano, N. R., Capaccioli, M., Romanowsky, A. J., Douglas, N. G., Merrifield, M. R., Kuijken, K., Arnaboldi, M., Gerhard, O., Freeman, K. C. 2005, MNRAS, 357, 691
Neistein, E., van den Bosch, F. C., & Dekel, A. 2006, MNRAS, 372, 933
Nelan J. E., Smith R. J., Hudson M. J., Wegner G. A., Lucey J. R., Moore S. A. W., Quinney S. J., & Suntzeff N. B. 2005, ApJ, 632, 137
Ogando R.L.C. et al. 2005, ApJ, 632, 61
Padmanabhan, N. et al. 2004, NewA, 9, 329P
Pannella M., Hopp U., Saglia R. P., Bender R., Drory N., Salvato M., Gabasch A., Feulner G., 2006, ApJl, 639, L1
Panter B., Jimenez R., Heavens A. F., Charlot S., 2007, MNRAS, 378, 1550
Park, C. & Choi, Y.-Y. 2005, ApJ, 635, 29
Peletier R.F., Davies R.L., Illingworth G.D., Davis L.E., & Cawson, M. 1990a, AJ, 100, 1091
Peletier R.F., Valentijn E.A., & Jameson, R.F. 1990b, A&A, 233, 62
Peletier R. F et al. 2007, MNRAS, 379, 445
Pipino A., Matteucci F., Borgani S., Biviano A., 2002, New Astron., 7, 227
Pipino A., D'Ercole A. & Matteucci F. 2008, A&A, 484, 679
Proctor R. N. 2003, PhD Thesis, University of Central Lancashire
Proctor R. N., Forbes D. A., Forestell A., Gebhardt K., 2005, MNRAS, 362, 857
Prugniel, Ph. & Simien F. 1997, A&A, 321, 111
Rawle T.D., Smith R.J. & Lucey J.R., 2009, arXiv:0909.3844
Recchi S., Matteucci F. & D'Ercole A., 2001, MNRAS, 322, 800
Reda F. M., Proctor R. N., Forbes D. A., Hau G. K. T. & Larsen, S. S. 2007, MNRAS, 377, 1772
Rettura A. et al. 2006, A&A, 458, 717
Romeo A. D., Napolitano N. R., Covone G., Sommer-Larsen J., Antonuccio-Delogu V., & Capaccioli M. 2008, MNRAS, 389, 13
Ruszkowski M. & Springel V. 2009, arXiv:0902.0373
Saglia, R.P., Maraston, C., Greggio, L., Bender, R., & Ziegler, B. 2000, A&A, 360, 911
Salpeter, E.E. 1955 ApJ, 121, 161
Sánchez-Blázquez P., Forbes D. A., Strader J., Brodie J., Proctor R. 2007, MNRAS, 377, 759
Sánchez-Blázquez P., Courty S., Gibson B. K. & Brook C. B. 2009, MNRAS 398, 591
Sandage, A. 1972, ApJ, 176, 21
Scannapieco C., Tissera P. B., White S. D. M., Springel V., 2006, MNRAS, 371, 1125
Schawinski K., Thomas D., Sarzi M., Maraston C., Kaviraj S., Joo S.-J., Yi S. K., Silk J. 2007, MNRAS, 382, 1415
Schlegel D. J., Finkbeiner D. P., Davis M., 1998, ApJ, 500, 525
Shen, S., Mo, H.J., White, S.D.M., Blanton, M.R., Kauffmann, G., Voges, W., Brinkmann, J., Csabai, I. 2003, MNRAS, 343, 978
Shimasaku K., Fukugita M., Doi M., Hamabe M. et al. 2001, AJ, 122, 1238
Sijacki D. et al. 2007, MNRAS, 380, 877
Spergel, D.N. et al. ApJS, 2003, 148, 175
Spergel, D. N., et al. 2007, ApJS, 170, 377
Spolaor, M., Proctor, R. N., Forbes, D. A., Couch, W. J. 2009, ApJ, 691, 138
Springel V., Di Matteo T. & Hernquist L. 2005, ApJ, 620, 79
Strateva I. et al., 2001, AJ, 122, 1861
Suh H., Jeong H., Oh K., Yi S. K., Ferreras I. & Schawinski K. 2010, ApJS, 187, 374
Tamura, N., et al. 2000, AJ, 119, 2134
Tamura, N., & Ohta, K. 2000, AJ, 120, 533
Tamura, N., & Ohta, K. 2003, AJ, 126, 596
Taylor V.A. et al. 2005, ApJ, 630, 784
Thomas D., Maraston C., Bender R. & Mendes de Oliveira C. 2005, ApJ, 621, 673
Tortora C. et al. 2009a, MNRAS, 396, 1132
Tortora C. et al. 2009b, MNRAS, 396, 61

Trager S. C., Faber, S. M., Worthey, G., Gonzlez, J. J. 2000, AJ, 119, 1645
Treu, T., et al. 2005, ApJ, 633, 174
Trujillo I. et al. 2004, ApJ, 604, 521
Trujillo, I. et al. 2006, ApJ, 650, 18
Tully R. B., Verheijen M. A. W., Pierce M. J., Huang J. & Wainscoat R. J. 1996, AJ, 112, 2471
Vader J. P., Vigroux L., Lachièze-Rey M., & Souviron, J. 1988, A&A, 203, 217
van der Wel et al. 2006, ApJ 652, 97
Weinmann S.M., Kauffmann G., van den Bosch F.C., Pasquali A., McIntosh D.H., Mo H., Yang X., Guo Y., 2009, MNRAS, 394, 1213
White S. D. M. 1980, MNRAS, 191, 1
Wise M. W. & Silva D. R. 1996, ApJ, 461, 155
Worthey G. 1994, ApJS, 95, 107
Wu H., Shao Z., Mo H. J., Xia X., Deng Z. 2005, ApJ, 622, 244

APPENDIX A: SAMPLE SYSTEMATICS AND CONTAMINANTS

We analyze various systematics in the galaxy sample and in the selection criteria.

A1 Sample systematics

We quantify the effect of source incompleteness on our results, in particular on the age and metallicity gradients as a function of stellar mass.

Our sample is magnitude limited, and misses increasingly fainter and less massive galaxies at larger redshifts; e.g. at $z_{spec} = 0.02$, the lowest mass galaxies have $M_* \sim 10^{8.2}\,M_\odot$, while at $z_{spec} = 0.05$ this limit increases to $M_* \sim 10^9\,M_\odot$ (see Blanton et al. 2005a and the left panel in Fig. A1). We already mentioned that dwarf galaxies can be missed due to surface brightness selection effects (photometric incompleteness). Other less obvious sources of incompleteness are defects in deblending and the absence of reliable redshift measurements.

By analyzing the surface brightness completeness of this sample of galaxies, Blanton et al. (2005a) found that for $r < -18$ mag ($M_* \lesssim 10^9\,M_\odot$) the catalog is complete at more than 90%, while at fainter magnitudes the photometric pipeline could mistakenly deblend galaxies or overestimate the background sky level. However, down to $M_* \sim 10^{8.5}\,M_\odot$ the overall incompleteness is not dramatic (Blanton et al. 2005a, Baldry, Glazebrook & Driver 2008, Li & White 2009). Moreover, at $z_{spec} \lesssim 0.02$ the distribution of galaxies seems patchy and the fraction of massive galaxies with $M_* \gtrsim 10^9\,M_\odot$ turns out to be smaller, and not homogeneously distributed as the one at larger redshifts. In order to account for all possible biases due to these missing galaxies we have redone the age and metallicity gradient plots for the subsamples of galaxies having $z_{spec} > 0.02$ and $z_{spec} > 0.03$, in order to avoid the regions with higher incompleteness. The trends found for the two subsamples are almost identical to the ones for the full data set, which suggests that the impact on our results of all sources of incompleteness is negligible, both at low and high masses.

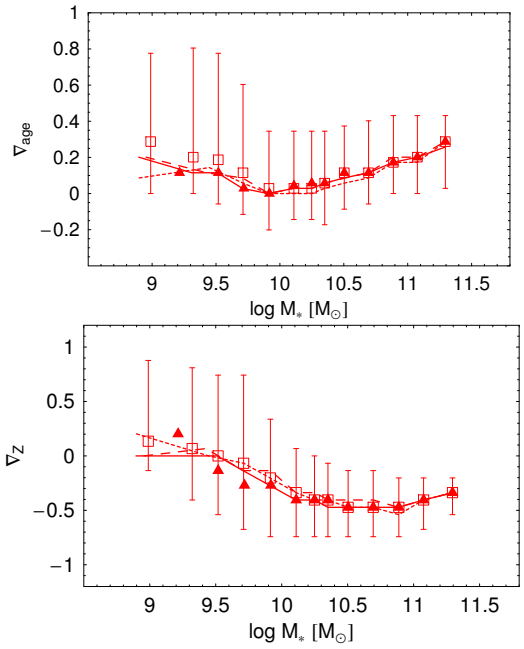

Figure A2. Systematics in the $\nabla_{\rm age}$ and $\nabla_{\rm Z}$ as a function of stellar mass due to changes in selection criteria. Boxes and bars are for our reference results. Continue, long-dashed, and short dashed lines are for a) $C > 2.6$, b) $C > 3$ and c) $2.5 < n < 5.5$ respectively, and triangles for our reference criteria with a cut in the total color ($g - r > 0.55$).

It has been suggested that the Sérsic fitting procedure adopted in B05 contain significant biases. In particular, at very high luminosities, Sérsic indices, effective radii and fluxes can be underestimated by ~ 0.5, $\sim 10 - 15\%$ and $\sim 10\%$, respectively (B05). Similar considerations hold for systems with high effective radii and Sérsic indices. The Sérsic procedure used in B05 have been extensively analyzed in Guo et al. (2009), who have shown that the main source of systematics is introduced by the 1D Sérsic fit and background sky level estimate. However, these systematics turn out to produce shallower CGs (of ~ 0.03) than our estimates, thus leaving the main trends with luminosity and mass unaffected. On the other hand, for the systems with very low Sérsic indices, the B05 fitting procedure recovers quite well the parameters, although a slight underestimate of fluxes and an overestimate of n might still occur (Guo et al. 2009).

A2 Morphological separation and contamination

Seeking for a trustful morphological classification, the eyeball check or the B/T are the best approaches, but unfortunately they are not suitable procedure for very large sample of galaxies (Allen et al. 2006, Lintott et al. 2008). Indeed, the use of some objective automated morphological classification is more efficient. For these reasons, we decided to adopt a classification with the minimal parameter set to minimize the probability of erroneously discarding ETGs, even if this would make the LTG contamination non-negligible. We have "a posteriori" checked the effect of such a contamination on the results. We have verified that the use of the *concentration* and Sérsic $n-$index

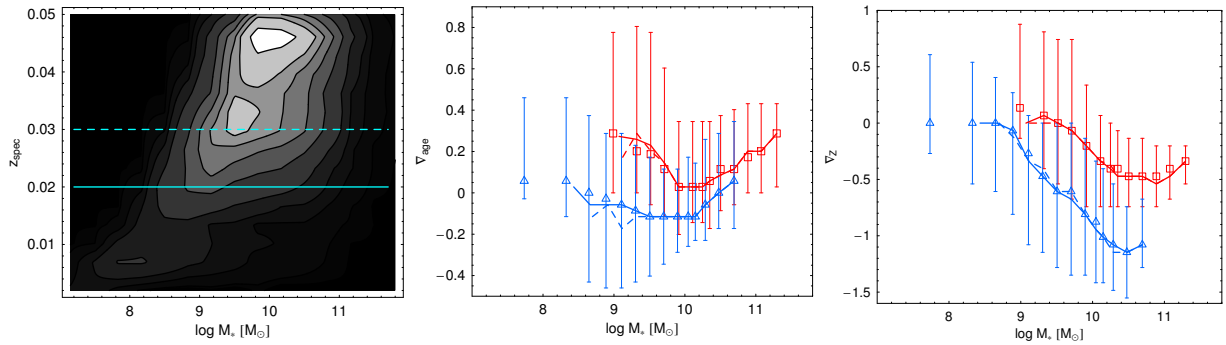

Figure A1. *Left Panel.* Spectroscopic redshift as a function of stellar mass for the B05 galaxy sample. The two lines are relative to the redshift thresholds we use to select subsamples of galaxies. *Middle Panel.* Age gradients as a function of stellar mass. *Right Panel.* Metallicity gradients as a function of stellar mass. The red and blue symbols with bars are relative to the full galaxy sample for ETGs and LTGs respectively, as discussed in the text, and the continue and dashed lines are for subsamples with $z_{spec} > 0.02$ and $z_{spec} > 0.03$.

parameters is efficient to this purpose. In fact, the Sérsic index is a measure of the steepness of the surface brightness profile and is historically considered as a good indicator of galaxy type: LTGs are well fitted by light profiles with lower n, being $n = 1$ the prototype of light profile for galaxy disks, and $n > 1$ the ones typical of ETGs. Recently, the concentration parameter determined from the SDSS pipeline has been claimed to be an even better proxy of galaxy morphology (Shimasaku et al. 2001, Strateva et al. 2001, Nakamura et al. 2003, Weinmann et al. 2009). We have checked that the results on the gradient trends do not strongly depend on the constraints adopted on the two parameters for the ETG/LTG separation (see Fig. A2) with some significant change raising for age gradients at $M_* \lesssim 10^{9.5}\, M_\odot$. We also analyzed the impact on our results when a cut in colour is applied.

As anticipated, our selection criteria are not meant to be contamination free, with ETGs being possibly contaminated by disk galaxies: e.g., both n and C depend on the inclination and edge-on disk or spiral galaxies can show parameter values typical of ETGs. An approach to avoid this kind of contamination can rely on the axial ratio, but in this case we would loose edge-on S0s misclassified as LTGs (Maller et al. 2009, Weinmann et al. 2009). In order to evaluate the fraction of contaminants in our ETG sample we have made an eyeball inspection of a subsample of 300 randomly selected galaxies. This classification may be still prone to subjectivity, thus four of us have checked the galaxies and produced an average classification. Eight galaxy images are shown in Fig. A3. We find that $\sim 70 - 75\%$ of galaxies are ETGs, while the remaining galaxies are uncertain objects or clear spirals. These uncertain or misclassified galaxies mainly fill the intermediate-mass region and many of them have red colors $g - i \gtrsim 0.5$. Although the fraction of such contaminants is not negligible, we have found that the gradients are not strongly biased on average, as shown in Fig. A4. The bottom line of this check is that the inclusion of contaminants leaves the average gradients unaffected. For instance, on the mass range $10 < \log M_* < 11$, we obtain as median value $\nabla_{\rm age} = 0.12$, and 0.03 for the ETGs, and the contaminants respectively, as shown in Fig. A4, while we find $\nabla_{\rm Z} = -0.47$ and -0.54 for the same systems. However, when computing the median values for the whole sample, the result is unchanged from the one of systems correctly classified as

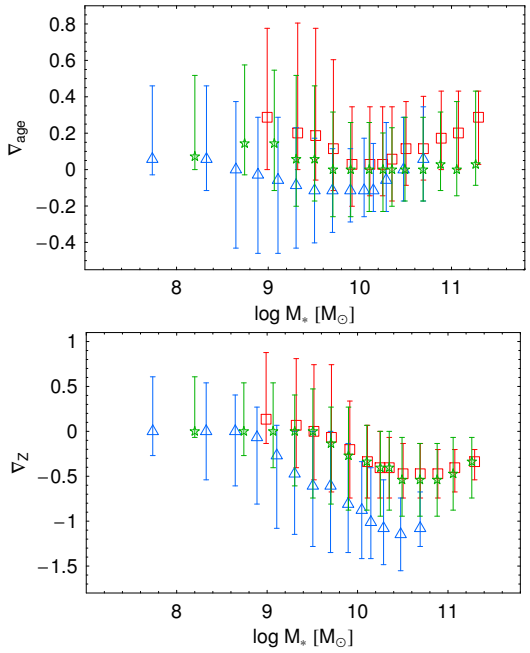

Figure A5. Age (left panel) and metallicity (right panel) gradients as a function of stellar mass for ETGs (red symbols), LTGs (blue symbols) and ITGs (green symbols).

ETGs. No net difference in the central galaxy age is found, while central metallicity for contaminants is larger than the one for ETGs, consistently with the steeper metallicity gradients for galaxies with larger metallicity (see Fig. 6). We have also checked that the presence of contaminants in ETG sample, in particular, does not affect the dependence on the central age of $\nabla_{\rm age}$ and $\nabla_{\rm Z}$ as discussed in Fig. 4. Thus, we can safely assume that the central age is a genuine parameter which rules the scatter of the relations with mass and σ_0.

Finally, the conservative choice of both n and C provided a quite sharp separation between ETGs and LTGs. For sake of completeness, the residual galaxy sample with a mixed morphology, i.e. ITGs, turned out to have also an intermediate distribution of the age gradients as shown in Fig. A5. On the contrary, metallicity gradients of ITGs have a trend which is more similar to the one of ETGs.

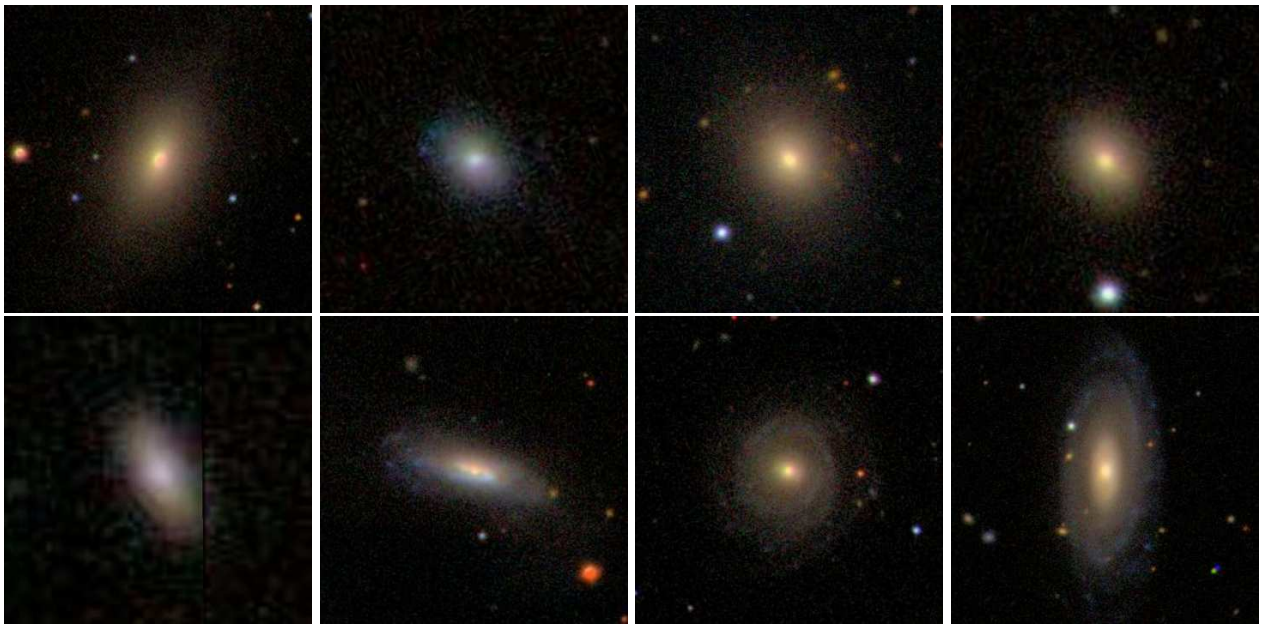

Figure A3. Images of 8 galaxies classified by eyeball checking. The galaxies in the top panels are classified as E/S0, the first one in the bottom panels is uncertain, while the last three are clear example of spirals.

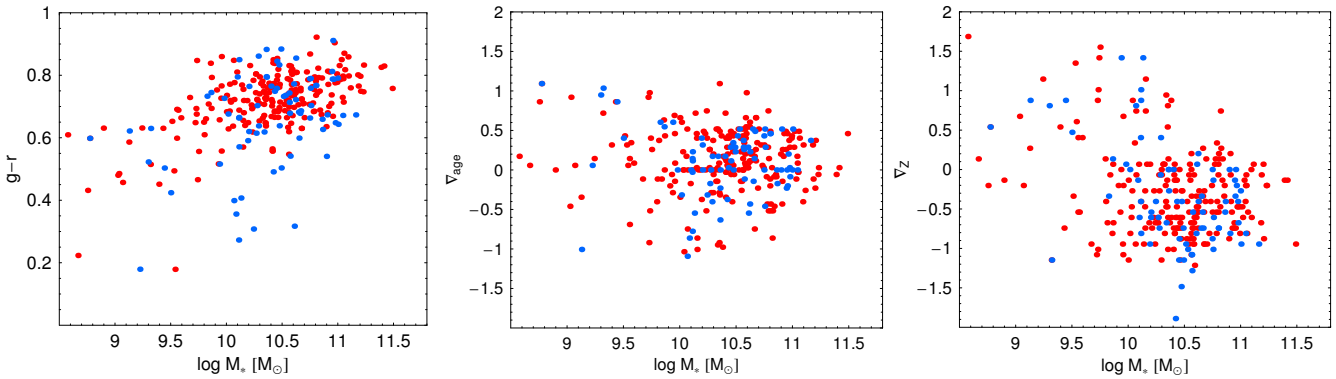

Figure A4. Results for galaxies extracted for eyeball check. The red points are for galaxies classified as ellipticals, while blue ones are for the galaxies with an uncertain classification or spirals. *Left Panel.* Color-mass diagram. *Middle Panel.* $\nabla_{\rm age}$ as a function of stellar mass. *Right Panel.* $\nabla_{\rm Z}$ as a function of stellar mass.

APPENDIX B: SYSTEMATICS IN STELLAR FITTING

We examine the role of systematic uncertainties on the stellar populations results. First, we address the problem of degeneracy in synthetic spectral models (§ B1), then derive the optical-infrared and infrared CGs from our fitted stellar populations and check the consistency with results from literature (§ B2). Dust extinction is discussed in § B3.

B1 Parameter degeneracies and wavelength coverage

We have analyzed the optical CGs, fitting the internal and external colors with synthetic spectral models to derive the age and metallicity of stellar populations. Because of the well known age-metallicity degeneracy (Worthey 1994, Bruzual & Charlot 2003, Gallazzi et al. 2005), the stellar parameters and the relative gradients might be severely biased, in particular by using the optical colors only. Increasing the wavelength coverage or adding spectral information would be the best approach to the problem.

Near-IR bands are primarily sensitive to the red stellar populations which represent the main fraction of stars forming an evolved galaxy, allowing to trace the visible mass of old galaxies, and are less affected by dust extinction. Their inclusion is suitable to alleviate the age-metallicity degeneracy (de Jong 1996, Cardiel et al. 2003, MacArthur et al. 2004, Chang et al. 2006), but is not expected to be definitive. Different colors are found to have a different sensitivity to age and metallicity (Li et al. 2007). For instance, both optical and infrared colors are found to be sensitive to metallicity, while age mostly affects the optical colors (Chang et al. 2006). Wu et al. (2005) showed that the inclusion of near-IR bands allows one to recover much more accurate stellar parameters from the fit, but the confidence contours are still

Table B1. Results from Montecarlo simulations. We show the median relative age $\Delta(age) = (age_{fit} - age_{in})/age_{in}$ and metallicity $\Delta(Z) = (Z_{fit} - Z_{in})/Z_{in}$, where *fit* and *in* are for estimated and input parameters. The results for different initial perturbations δ and without or with IR photometry are shown.

		ugriz	ugrizJHKs
$\Delta(age)$	$\delta = 0.01$	$0.01^{+0.12}_{-0.10}$	$0^{+0.08}_{-0.07}$
	$\delta = 0.03$	$0.02^{+0.27}_{-0.20}$	$0^{+0.12}_{-0.14}$
	$\delta = 0.05$	$0.03^{+0.49}_{-0.26}$	$0^{+0.25}_{-0.17}$
$\Delta(Z)$	$\delta = 0.01$	$-0.01^{+0.13}_{-0.13}$	$0^{+0.07}_{-0.08}$
	$\delta = 0.03$	$0^{+0.28}_{-0.27}$	$0^{+0.13}_{-0.12}$
	$\delta = 0.05$	$-0.02^{+0.39}_{-0.40}$	$0^{+0.19}_{-0.20}$

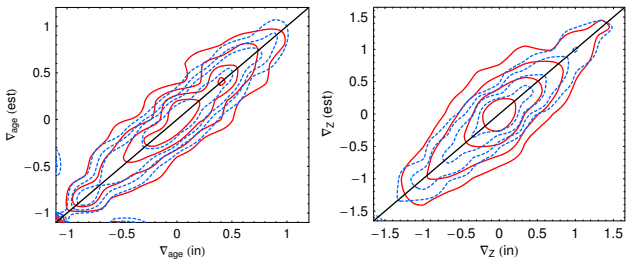

Figure B1. Comparison of age (left) and metallicity (right) gradients estimated from Montecarlo repetitions (y-axis), given the initial input values (x-axis). Here we consider our extreme error budget, $\delta = 0.05$. Red and dashed blue contours are the isodensity corresponding to the results from the stellar population fit, using only optical and optical+infrared colors respectively.

affected by the age-metallicity relation, and the gain with respect the optical bands is minimal. It is reasonable to argue that the use of a large statistical sample, as we adopt in this paper, is more effective in reducing the overall uncertainties of the galaxy properties than the adoption of a larger filter baseline.

Of course, the adoption of spectroscopic line absorption indices would be the best indicators of stellar population parameters and gradients, as they permit to break the age-metallicity degeneracy (e.g., Burstein et al. 1984, Henry & Worthey 1999, Carollo et al. 1993, Davies et al. 1993, Mehlert et al. 2003). However, spectral indices of large galaxy sample are still difficult to obtain for the observational limits imposed by the low starlight fluxes (e.g. around $R_{\rm eff}$ and beyond). Recently, integral-field spectroscopy has been used to determine gradients in line absorption indices (e.g. Kuntschner et al. 2006, Rawle et al. 2009) of elliptical and lenticular galaxies, but this approach is only applicable to limited samples of systems.

Despite the limited wavelength coverage, but taking advantage of the large statistical sample we have shown that both spectroscopy and photometric analysis converge to similar results for stellar population gradients (see e.g. §3.4). Similarly, Rawle et al. (2009) have shown that optical colors and CGs derived from spectroscopic age and metallicity are generally consistent with values directly recovered from photometry. Moreover, the recovered correlations between metallicity and age gradients are not spurious and not generated by correlated measurement errors on these quantities.

However, in the attempt of putting our result further on a more solid ground, we have tested the reliability of our color modelling technique and checked for the presence of spuriously generated correlations between age and Z or gradients $\nabla_{\rm age}$ and $\nabla_{\rm Z}$, running several Monte Carlo simulations. We extracted 1000 simulated galaxy spectra from our BC03 SED libraries with random (age_i, Z_i) for $i = 1, 2$ (i.e. with no correlations among these parameters and with $\nabla_{\rm Z}$, and $\nabla_{\rm age}$ randomly distributed), and applied our fitting procedure, comparing the estimated parameters with the input model values. We perform the fit 1) only using the optical SDSS bands ugriz and 2) adding the JHKs photometry from 2MASS survey (Jarrett et al. 2003), to understand the systematics that the fit of optical colors can induce. Typical observational uncertainties for each band are assigned (i.e. $\delta u \sim 0.06$, $\delta g \sim 0.035$, $\delta r \sim 0.035$, $\delta i \sim 0.035$, $\delta g \sim 0.04$, $\delta J \sim 0.08$, $\delta H \sim 0.12$, $\delta Ks \sim 0.09$, from matched catalog in B05) and the input model colors are perturbed adding a random step in the interval $(-\delta, +\delta)$ with $\delta = 0.01, 0.03, 0.05$ to account for measurement systematics. Results of this analysis are collected in Table B1 and Fig. B1. We found that, on average, both the stellar parameters and gradients are fairly well recovered from the fitting procedure, with a scatter increasing with δ, and no spurious correlations are induced at more than 99% confidence level (see also Wu et al. (2005) and Tortora et al. (2009a) for similar analysis). We confirm the finding of Wu et al. (2005) that the inclusion of near-IR bands allows a better recovery of the input stellar parameters with a scatter in the range $5 - 25\%$, while the scatter is larger (in the range $10 - 50\%$) when only SDSS bands are used. Thus, the most remarkable result of this analysis, is that, notwithstanding the larger scatter, we are able to successfully recover the stellar population parameters, even without near-IR photometry.

B2 Comparison with IR CGs

One of the advantages of synthetic spectral modelling is to allow the recover of synthetic IR colors by using the best fitted spectral parameters. Therefore, we derive the colors $g-J$, $g-H$ and $g-Ks$ at $0.1R_{\rm eff}$ and $1R_{\rm eff}$ and the relative CGs, which show trends with mass similar to those discussed in this paper. If we select luminous ($r < -20$ mag) ETGs we find median values of $\nabla_{g-J} = -0.33$, $\nabla_{g-H} = -0.37$ and $\nabla_{g-Ks} = -0.40$, with 25-75 percentiles in sample distribution of 0.17, 0.18, and 0.20, respectively. As already extensively discussed in the text, to compare these synthetic CGs with the observed values in La Barbera et al. (2009), we select older systems with $age_1 > 6$ Gyr. In this case we have $\nabla_{g-J} = -0.26$, $\nabla_{g-H} = -0.29$, and $\nabla_{g-Ks} = -0.31$ (with scatter of 0.16, 0.18, and 0.20), which are fairly consistent with the results reported by these authors. A fairly good agreement is also found with the median $\nabla_{g-Ks} = -0.29 \pm 0.07$ in Wu et al. (2005).

B3 Internal dust extinction

In our analysis we have assumed no internal extinction. This assumption is more reasonable for ETGs, where only a negligible fraction of interstellar matter is in the form of dust (e.g. Kormendy & Djorgovski 1989, Wise & Silva 1996, Dorman et al. 2003), but it might be unappropriate for LTGs, as these systems have a larger amount

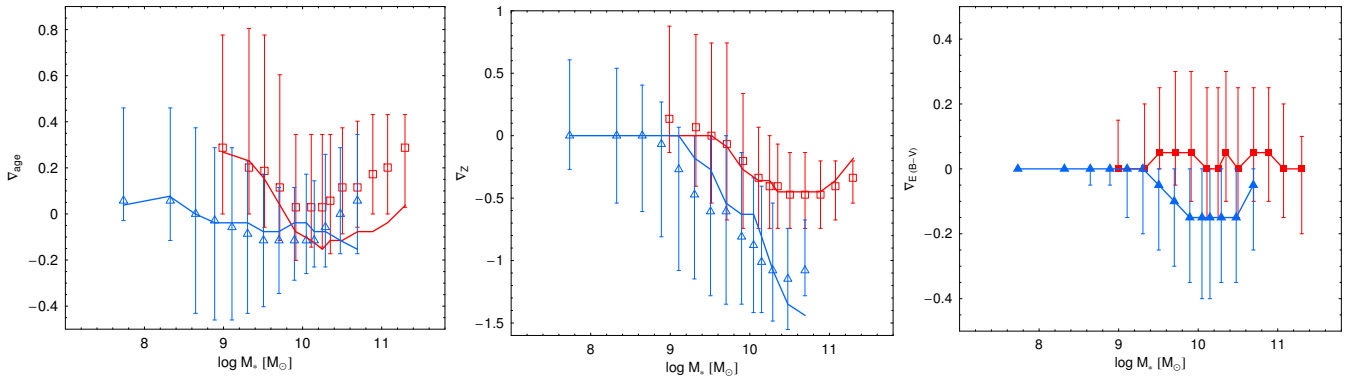

Figure B2. Gradients for the spectral model leaving free the galaxy age, metallicity, and dust content. The open symbols with bars are for our reference fit, while the line is for the fit with dust extinction free to vary. *Left Panel.* $\nabla_{\rm age}$ as a function of stellar mass. *Middle Panel.* $\nabla_{\rm Z}$ as a function of stellar mass. *Right Panel.* Gradient in dust extinction as a function of stellar mass.

of internal dust (e.g. Brammer et al. 2009). Usually no dust extinction is included in the analysis of ETGs (Tamura et al. 2000, Tamura & Ohta 2000, Tamura & Ohta 2003, Wu et al. 2005), but dust gradients would have effects in CGs (e.g. Goudfrooij & de Jong 1995). However, for ETGs it has been shown that metal absorption lines gradients, being less affected by dust extinction than colors (e.g. MacArthur 2005), are unlikely to be produced by pure dust extinction gradients to mimic the population gradients inferred by broadband colors (Peletier et al. 1990a, Davies et al. 1993, Carollo et al. 1993, Kobayashi & Arimoto 1999, Rawle et al. 2009). Thus they are not dominant but could still contribute to a fraction of CGs. Similarly, in LTGs the dust seems to have a negligible effect in shaping CGs (e.g. MacArthur et al. 2004), even thought there are still open questions (MacArthur et al. 2009) which might lead to a different conclusion. To analyze the impact of dust extinction on our results, we fit the colors at $0.1R_{\rm eff}$ and $1R_{\rm eff}$ using a library of spectra with age, metallicity and dust content free to change. The extinction curve of Cardelli, Clayton & Mathis (1989) is used to account for the extinction contribution in different wavebands and the color excess $E(B-V)$ in the range $0-0.5$ is used as a free parameter. The results are shown in Fig. B2, where they are compared with our reference results with no dust. The main trends are unaffected by the dust, although we note some differences in the $\nabla_{\rm age}$ at $M_* \gtrsim 10^{9.5}\,M_\odot$. On the other hand, only slight differences are observed for metallicity gradients, and the qualitative trends are unaffected both in ETGs and LTGs. Thus this analysis has shown that, although dust extinction is not negligible in both ETGs and LTGs, the trends and the amplitude of gradients are still dominated by changes in stellar population parameters. We show the dust extinction gradients, defined as $\nabla_{E(B-V)} = E(B-V)_2 - E(B-V)_1$ in the right panel of Fig. B2. On average, the ETGs exhibit null or slightly positive dust gradients (with a tail towards negative values, too), while LTGs have opposite behaviour with a larger dust content in the inner regions, showing steeper dust gradients at intermediate high masses (e.g., Dalcanton & Bernstein 2002, MacArthur et al. 2004). Finally, we have checked the absence of any correlation between age and metallicity gradient with both $\nabla_{E(B-V)}$ and the central extinction for LTGs which has suggested that the results of these systems are not affected by the presence of dust. We remark that the use of only optical bands and the addition of IR photometry is not sufficient to break the degeneracy between stellar population parameters and dust, for this reason we are tempted to consider these results more as a caveat for this particular data-set rather than a conclusive check. If the correction of the age gradients for the dust extinction is right, the net effect is that the age gradient might be less affecting the color gradient for massive systems, and only metallicity gradient should matter.